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A Trio of GRB-SNe:

GRB 120729A, GRB 130215A / SN 2013ez and GRB 130831A / SN 2013fu

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ABSTRACT

We present optical and near-infrared (NIR) photometry for three gamma-ray burst supernovae (GRB-SNe): GRB 120729A, GRB 130215A / SN 2013ez and GRB 130831A / SN 2013fu. In the case of GRB 130215A / SN 2013ez, we also present optical spectroscopy at $t - t_0 = 16.1$ d, which covers rest-frame 3000–6250 Å. Based on Fe II $\lambda 5169$ and Si II $\lambda 6355$, our spectrum indicates an unusually low expansion velocity of ~ 4000 – 6350 km s⁻¹, the lowest ever measured for a GRB-SN. Additionally, we determined the brightness and shape of each accompanying SN relative to a template supernova (SN 1998bw), which were used to estimate the amount of nickel produced via nucleosynthesis during each explosion. We find that our derived nickel masses are typical of other GRB-SNe, and greater than those of SNe Ibc that are not associated with GRBs. For GRB 130831A / SN 2013fu, we use our well-sampled *R*-band light curve (LC) to estimate the amount of ejecta mass and the kinetic energy of the SN, finding that these too are similar to other GRB-SNe. For GRB 130215A, we take advantage of contemporaneous optical/NIR observations to construct an optical/NIR bolometric LC of the afterglow. We fit the bolometric LC with the millisecond magnetar model of Zhang & Mészáros (2001), which considers dipole radiation as a source of energy injection to the forward shock powering the optical/NIR afterglow. Using this model we derive an initial spin period of $P = 12$ ms and a magnetic field of $B = 1.1 \times 10^{15}$ G, which are commensurate with those found for proposed magnetar central engines of other long-duration GRBs.

Key words. words. that are about keys.

1. Introduction

Observational evidence supporting the connection between long-duration gamma-ray bursts (GRBs) and stripped-envelope, core-collapse supernovae (SNe) is now quite extensive (see Woosley & Bloom, 2006, and Hjorth & Bloom 2012 for extensive reviews of gamma-ray burst supernovae; GRB-SNe). 2013 was a prosperous year for GRB-SN science, with no less than four spectroscopic GRB-SN associations: GRB 130215A / SN 2013ez (de Ugarte Postigo et al. 2013b); GRB 130427A / SN 2013cq (de Ugarte Postigo et al. 2013c; Xu et al. 2013a; Levan et al. 2013; Melandri et al. 2014); GRB 130702A / SN 2013dx (Schulze et al. 2013) and GRB 130831A / SN 2013fu (Klose et al. 2013; Nicuesa Guelbenzu et al. 2013). These events join other spectroscopic GRB-SN associations (e.g. Galama et al. 1998; Hjorth et al. 2003; Stanek et al. 2003; Della Valle et al. 2003; Malesani et al. 2004; Pian et al. 2006; Chornock et al. 2010; Bufano et al. 2012; Berger et al. 2011; Sparre et al. 2011; Klose et al. 2012; de Ugarte Postigo et al. 2012; Melandri et al. 2012; Jin et al. 2013; Schulze et al. 2014). Numerous photometric inferences of GRB-SNe via SN bumps in optical and

near-infrared (NIR) light-curves (LCs; e.g. Zeh et al. 2004) further strengthen the GRB-SN connection (see Cano 2013 for a review).

The favoured physical description for producing a GRB is the “collapsar” scenario (Woosley 1993; MacFadyen & Woosley 1999; MacFadyen et al. 2001), where a compact object forms during the collapse of a massive star and ejects shells of material collimated in a jet at relativistic velocities. Multiple shells interact producing the initial γ -ray pulse, and as they propagate away from the explosion they encounter circumstellar material (CSM) ejected by the progenitor star prior to explosion (as well as interstellar material), producing a long-lived afterglow (AG). In the simplest scenario, a forward shock (FS) is thought to be created when the shells interact with the CSM, which accelerate electrons that cool by emitting synchrotron radiation. A couple of weeks (rest-frame) after the initial γ -ray pulse energetic SNe are then observed at optical and NIR wavelengths.

A basic assertion of the collapsar model is that the duration of the GRB prompt phase is the difference between the time that the central engine operates (i.e. T_{90} ; though see Zhang et al. (2014) who argue that T_{90} is not a reliable indicator of the engine activity timescale) minus the time it takes for the jet to breakout

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of the star: $T_{90} \sim t_{\text{engine}} - t_{\text{breakout}}$. A direct consequence of this premise is that there should be a plateau in the distribution of T_{90} for GRBs produced by collapsars when $T_{90} < t_{\text{breakout}}$, which was confirmed by Bromberg et al. (2012). Moreover, the value of T_{90} found at the upper-limit of the plateau seen in three satellites (BATSE, *Swift* and FERMI) was approximately the same ($T_{90} \sim 20\text{--}30$ s), which is the typical breakout time of the jet. This short breakout time suggests that the progenitor star at the time of explosion is very compact ($\sim 5 R_{\odot}$; Piran et al. 2013). Bromberg et al. (2013) then used these distributions to calculate the probability that a given GRB arises from a collapsar or not based on its T_{90} and hardness ratio.

The theoretical and observational evidence for the GRB-SNe connection is strong, however some questions remain unanswered. One of the biggest uncertainties is the nature of the compact object that powers the GRB. One possible engine is a stellar black hole rapidly accreting mass from a torus (e.g. Woosley 1993; MacFadyen et al. 2001). Another scenario is the extraction of energy via the spin-down of a compact object through magnetic winds, either a neutron star with a very large magnetic field ($10^{15\text{--}16}$ Gauss) and rotating near breakup ($P \approx 1$ ms; i.e. a millisecond magnetar), or through the spin-down of a rapidly rotating (i.e. Kerr) black hole (e.g. van Putten et al. 2011). Numerous flares and plateaus have been seen in AG LCs at X-ray and optical wavelengths (e.g. O’Brien et al. 2006, Margutti et al. 2010; Grupe et al. 2007, 2010), which require energy injection from a central engine. The origin of the energy injection is still uncertain however, and may arise from different sources in different events.

Secondly, the nature of the progenitor system has yet to be determined. Due to the vast cosmological distances at which GRBs occur it is generally not possible to detect the progenitor directly, as has been done for the progenitors of other types of core-collapse SNe (e.g. Smartt et al. 2009; Maund et al. 2014). Instead, the possible configuration of the progenitor system has to be indirectly inferred, whereby it is a formidable challenge to resolve the ambiguity between single and binary stars. Arguments based on statistically significant sample sizes of the bolometric properties of GRB-SNe in relation to the other SN Ibc subtypes (Ib, Ic and Ic-BL; Cano 2013) indicate that the progenitors of most SNe Ibc likely arise from binary systems, where the mass of individual stars in the system is less than what is attributed to single Wolf-Rayet stars observed in nature (Crowther et al. 2007). In these systems the outer layers of the star are tidally stripped, as well as ejected via line-driven winds.

Conversely, the progenitors of SNe Ic-BL and GRB-SNe may arise from more massive single-star progenitors (e.g. Yoon & Langer 2005; Woosley & Heger 2006, Yoon et al. 2012), where rapid rotation mixes the interior of the star. In these models a consequence of the rapid rotation is that the stars avoid a red supergiant stage, and there is a lack of strong coupling between the stellar core and surrounding envelope. This allows more angular momentum to be retained in the core, and any mass loss experienced by the star (which also removes angular momentum) is reduced as it is lost from a much smaller surface area. An expectation of these models is that the progenitors of SNe Ic-BL are more metal rich than those of GRB-SNe, (which has also been observed in nature, e.g. Modjaz et al. 2011; though see as well Levesque et al. 2012, Krühler et al. 2012, Savaglio et al. 2012 and Elliott et al. 2013 who have shown, respectively, that GRBs 020819A, 080605, 090323 and 110918A occurred in galaxies of solar and super-solar metallicities), and therefore lose more mass before exploding than GRB-SNe. This provides a natural explanation for why a high-energy transient is observed

in the latter because the central engine that is formed has retained more angular momentum at the time of explosion (e.g. Woosley & Zhang 2007). However, GRBs may also arise via binary systems, where the system may undergo a common-envelope phase. If the system remains intact after one of the stars explodes, the in-spiral of the compact object into the core of the unexploded secondary can impart angular momentum to the core, which may be retained at the time of explosion to then power a GRB (e.g. Cantiello et al. 2007).

In this paper we attempt to address at least one of these outstanding questions, namely the nature of the compact object acting as a central engine of GRB 130215A. Using the model of Zhang & Mészáros (2001) we show that energy injection from a millisecond magnetar provides a plausible fit to an optical/NIR bolometric LC of the AG. Using simple assumptions of the magnetar’s mass and radius we derive physically plausible estimates of its magnetic field strength and initial spin period, finding they are consistent with those found for proposed magnetar central engines of other long-duration GRBs.

The other main focus of this work is an investigation of the observational and physical properties of three GRB-SNe. A key result we find is that SN 2013ez, associated with GRB 130215A, is more likely to be of type Ic than type IcBL, which all other GRB-SNe apart from SN 2002lt (associated with GRB 021211, Della Valle et al. 2003; see Section 5) have been spectroscopically classified as. If our interpretation of the absorption features near ~ 5100 and ~ 6200 Å as blueshifted Fe II $\lambda 5169$ and Si II $\lambda 6355$ (indicating blueshifted velocities of ≈ 4000 and ≈ 6350 km s $^{-1}$) is correct, it implies that SN 2013ez has the slowest ejecta velocities ever measured for a GRB-SN. In sections 2, 3 and 4 we present photometric observations of GRB 120729A, GRB 130215A / SN 2013ez and GRB 130831A / SN 2013fu¹, respectively. A SN signature is seen in each event, which arises via SN-bumps for GRBs 120729A and 130831A, and a bump+spectrum for GRB 130215A. In section 5 we discuss the observational and physical properties of these three GRB-SNe in relation to other SNe Ibc.

Throughout this paper we use a Λ CDM cosmology constrained by Planck (Planck Collaboration et al. 2013) of $H_0 = 67.3$ km s $^{-1}$ Mpc $^{-1}$, $\Omega_M = 0.315$, $\Omega_{\Lambda} = 0.685$. Foreground extinction has been calculated using the dust extinction maps of Schlegel et al. (1998) and Schlafly & Finkbeiner (2011), while values of the rest-frame extinction that have been derived from our data are presented in Table 1. All bolometric properties (nickel mass, ejecta mass and kinetic energy, M_{Ni} , M_{ej} and E_K , respectively) are calculated for the rest-frame filter range *UBVRIJH* using the method in Cano (2013; C13 hereafter). Unless stated otherwise, errors are statistical only. Observer-frame times are used unless specified otherwise in the text. The respective decay and energy spectral indices α and β are defined by $f_{\nu} \propto (t - t_0)^{-\alpha} \nu^{-\beta}$, where t_0 is the time at which the GRB triggered the BAT instrument on-board the *Swift* satellite.

2. GRB 120729A

GRB 120729A was detected at 10:56:14 UT on 29-July-2012 by the *Swift* Burst Alert Telescope (BAT), and has a $T_{90} = 71.5 \pm 17.5$ s in the 15–350 keV energy range (Ukwatta et al. 2012; Palmer et al. 2012). It was also detected by the *Fermi* Gamma-Ray Burst Monitor (GBM) with a $T_{90} \approx 25$ s in the

¹ SN 2013fu was spectroscopically associated with GRB 130831A by Klose et al. 2013, in this work we present optical photometry only.

Table 1. GRB-SNe: Vital Statistics

GRB	SN	z	$A_{V,\text{fore}}$ (mag)	$A_{V,\text{rest}}$ (mag)	d_L^\dagger (Mpc)
120729A	-	0.80	0.55	0.15	4910.7
130215A	2013ez	0.597	0.53	0.0	3453.5
130831A	2013fu	0.479	0.15	0.0	2664.2

[†] Luminosity distance calculated using $H_0 = 67.3 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.315$, $\Omega_\Lambda = 0.685$.

50–300 keV energy range (Rau 2012). Rapid follow-up by several ground-based telescopes identified an optical transient coincident with the XRT position (Virgili et al. 2012; Oates & Ukwatta, 2012; Im & Hong, 2012; Wren et al. 2012; Gorosabel et al. 2012; D’Avanzo et al. 2012), and a redshift of $z = 0.80$ was measured with Gemini-North (Tanvir & Ball, 2012). The AG was not detected at radio (Laskar et al. 2012) down to 3σ upper limits of 39 mJy and 58 mJy, at 5.8 and 21.8 GHz, respectively. The AG was also not detected at sub-mm wavelengths (350 GHz, with rms of 1.8 mJy; Smith et al. 2012). An estimate of the isotropic energy release in γ -rays ($1\text{--}10^4$ keV, rest-frame) is $E_{\text{iso},\gamma} = 2.3^{+0.3}_{-1.5} \times 10^{52}$ erg, while the peak energy $E_p \approx 310.6 \text{ keV}^2$. The probability that GRB 120729A arises from a collapsar (Bromberg et al. 2013) based on T_{90} alone is $99.996 \pm 0.001\%$ (BAT) and $98.225 \pm 1.004\%$ (GBM). We use a foreground extinction value of $E(B - V)_{\text{fore}} = 0.164$ mag for GRB 120729A.

2.1. Data Reduction & Photometry

We obtained observations with the 2-m Faulkes Telescope North (FTN) robotic telescope starting less than ten minutes after the γ -ray detection. Subsequent follow-up observations were obtained with the 2-m Liverpool Telescope (LT), the 0.82-m Instituto de Astrofísica de Canarias (IAC) IAC80 telescope, the 3.6-m Telescopio Nazionale Galileo (TNG), and the 10.4-m Gran Telescopio Canarias (GTC) telescope. Six epochs of GTC images in *griz* were obtained during the first month, and a final epoch in all filters at $t - t_0 \approx 190$ d that were used as templates for image subtraction. Image reduction of data obtained on all telescopes was performed using standard techniques in IRAF³.

Calibration of the GTC data has been performed using standard star photometry. Observations (in *griz*) of Landolt standard field PG1323-086 (Landolt 1992) were obtained the same night as the final GTC epoch, all of which were taken under photometric conditions. The BVR_cI_c magnitudes of PG1323-086 were transformed into *griz* using transformation equations from Jordi et al. (2006), and the subsequent calibration was done using a zeropoint between the instrumental and catalog magnitudes. The calibration in each filter was then used to create a set of secondary standards in the GRB field, which the GTC images are calibrated against.

The $gVR_c i$ FTN, LT and TNG images are shallower than the GTC ones, where common stars are either saturated in the GTC images or not visible in the FTN/LT ones. Instead these images have been independently calibrated via standard star photometry using Landolt standards taken with the TNG the same night as the GRB observations. A zeropoint was computed between

the Landolt standards and the instrumental $gVR_c i$ magnitudes, which was then used to create a small set of secondary standards in the GRB field that were visible in all FTN/LT/TNG images but not saturated. The R_c magnitudes were then transformed into r magnitudes using transformation equations from Lupton (2005)⁴, which requires colours between filters (e.g. $R_c - i$) in order to properly calculate the corresponding SDSS magnitudes. Early observations in R_c and i were taken within 5–10 minutes of each other, however, as these observations were taken very soon after the initial GRB trigger (after only a few tens of minutes), it may not be appropriate to assume no colour/spectral evolution during the timing of a given R_c -band and i -band observation. Instead we estimated magnitudes in r at the time of the R_c -band observation by two methods: (1) a log-linear interpolation of the i -band LC, and (2) fitting a broken PL to the i -band LC in order to extrapolate to earlier and later times. Once the fitted functions were determined, the i magnitude at the time of the R_c observation was outputted. When possible we have taken the average magnitude found with both of these methods, and included their standard deviation when calculating the SDSS magnitude in r .

We used our deep GTC images to obtain image-subtracted magnitudes of the optical transient (OT) associated with GRB 120729A, using the final epoch in each filter as a template. This method proved to be valuable to isolate the OT flux as the field is quite crowded, and because the OT is very faint. Already at $t - t_0 = 0.75$ d the *griz* magnitudes of the OT are $m_{AB} = 24\text{--}25$. Image subtraction was performed using an adaptation of the original ISIS program (Alard & Lupton 1998; Alard 2000) that was developed for Hubble Space Telescope SN surveys by Strolger et al. (2004). A key advantage of this code is the option for the user to specify a set of stamps for the program to use when it calculates the point-spread function in each image. The image-subtraction technique was then optimised by varying the kernel mesh size and measuring the standard deviation (σ) of the background counts in a nearby region in the image (where images with lower σ values indicate they are a better subtracted image). As a self-consistency check, we compared the OT magnitudes against those found by performing photometry on the un-subtracted images, converting the magnitudes into fluxes, and then mathematically subtracting the host flux away. Good agreement was obtained with both methods, showing that the image-subtraction technique was well optimised.

The *griz* magnitudes of the host galaxy were measured, and these magnitudes were converted into monochromatic fluxes using the flux zeropoints from Fukugita et al. (1995) and subtracted from the earlier observations obtained with the LT, FTN and IAC80. The apparent magnitudes (not corrected for foreground or host extinction) of the GRB+SN+host are presented in Table 3.

² http://butler.lab.asu.edu/swift/bat_spec_table.html

³ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

⁴ <http://www.sdss.org/dr4/algorithms/sdssUBVRITransform.html>

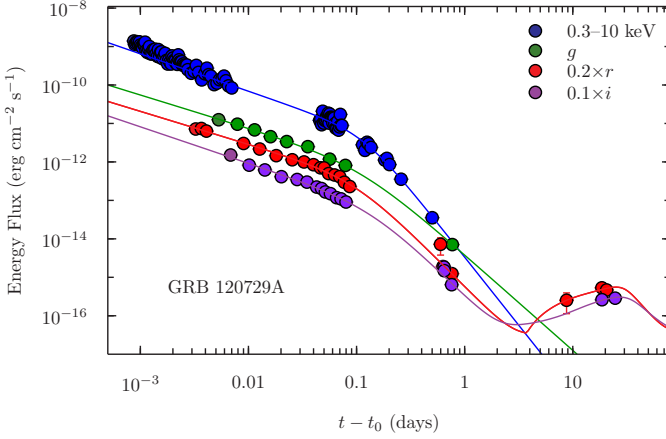


Fig. 1. GRB 120729A: Optical and X-ray (0.3–10 keV) light curves. $BR_c I_c$ magnitudes have been transformed into gri using transformation equations from Jordi et al. (2006), see the main text for details. The optical data are host-subtracted and have been corrected for foreground and rest-frame extinction. Note that the errorbars are usually smaller than the size of the plotted symbols. All LCs have been fit with a broken power-law in order to determine the decay rate before ($\alpha_{v,1}$) and after ($\alpha_{v,2}$) and the break ($T_{v,B}$), as well as the timing of the break. It is seen that α_1 is approximately the same in the optical and X-ray, as well as the time of the break ($T_B \approx 0.1$ d). After the break the X-ray decays at a faster rate than the optical filters. In r and i we simultaneously fit a SN-component (i.e. a stretch and luminosity factor relative to a redshifted, k-corrected template LC). The paucity of optical points limits our analysis, however when fixing the stretch factor to $s = 1.0$ in both filters, we find luminosity factors of $k_r = 1.29 \pm 0.19$ and $k_i = 0.76 \pm 0.11$.

2.2. The Afterglow

We combined our optical detections with the *Swift* XRT (0.3–10 keV) observations (Fig. 1), where the host-subtracted magnitudes are corrected for foreground and rest-frame extinction, and been converted into monochromatic fluxes (mJy), and then into energy fluxes ($\text{erg cm}^{-2} \text{s}^{-1}$) using the zeropoints and filter effective wavelengths from Fukugita et al. (1995).

We fit all LCs with a broken power-law (PL) (we also included a SN-component that is simultaneously determined when fitting the r and i LCs, see section 2.5) in order to determine the decay rate before ($\alpha_{v,1}$) and after ($\alpha_{v,2}$) the break, as well as the timing of the break ($T_{v,B}$). Our best-fitting parameters (fit between 0.005–30 d) are: (1) X-ray: $\alpha_{X,1} = 0.97 \pm 0.06$, $\alpha_{X,2} = 3.54 \pm 0.27$, $T_{X,B} = 0.12 \pm 0.02$ d; (2) optical: $\alpha_{g,1} = 0.86 \pm 0.04$, $\alpha_{g,2} = 2.49 \pm 0.14$, $T_{g,B} = 0.11 \pm 0.02$ d; $\alpha_{r,1} = 0.86 \pm 0.03$, $\alpha_{r,2} = 2.85 \pm 0.10$, $T_{r,B} = 0.10 \pm 0.02$ d; $\alpha_{i,1} = 0.95 \pm 0.07$, $\alpha_{i,2} = 2.77 \pm 0.22$, $T_{i,B} = 0.14 \pm 0.04$ d. The time the LC breaks is approximately the same at all frequencies ($T_B \approx 0.11$ d). The value of α_1 is roughly the same at all wavelengths before the break, and while α_2 is steeper in the X-ray than the optical, it is quite similar in all optical bands, though there is a hint that it is decay slightly slower in g , though of course the paucity of observations limits how much we can comment on this. If instead we fit the optical LCs simultaneously, assuming that the time the LC breaks and the decay constants before and after the break are the same, we find $\alpha_{\text{opt},1} = 0.89 \pm 0.09$, $\alpha_{\text{opt},2} = 2.70 \pm 0.18$, $T_{\text{opt},B} = 0.10 \pm 0.04$ d.

If the achromatic break at $t - t_0 \approx 0.11$ d is interpreted as a jet break, it is possible to estimate the angular width of the jet using equation 4 in Piran (2004). Assuming a density of $n = 1 \text{ cm}^{-3}$, and an isotropic kinetic energy in the ejecta $\equiv \eta E_{\text{iso},\gamma}$, where η is the radiative efficiency and we assume a value of $\eta = 0.2$, we estimate an opening angle of $\theta \approx 4.4^\circ$. Using $E_{\text{iso},\gamma} = 2.3^{+0.3}_{-1.5} \times 10^{52}$, this in turn implies a beaming-corrected γ -ray energy release of $E_{\theta\gamma} = (\frac{\theta^2}{2}) E_{\text{iso},\gamma} \sim 6.8 \times 10^{49}$ erg. If the density is higher, $n = 10$, the opening angle is larger ($\theta \approx 5.7^\circ$), and so is the beam-corrected kinetic energy ($E_{\theta,\gamma} \sim 1.2 \times 10^{50}$ erg).

2.3. The Spectral Energy Distribution

We combined our host-subtracted GTC magnitudes at $t - t_0 = 0.75$ d, which were corrected for foreground extinction and converted into monochromatic fluxes, with contemporaneous X-ray observations to construct an X-ray to optical spectral energy distribution (SED), with the intention of getting an estimate of the amount of rest-frame extinction (Fig. 2). We used the general procedure outlined in Guidorzi et al. (2009) when constructing the energy spectrum. As there are fewer X-ray photons at late times (the final observation is at $t - t_0 = 0.5$ d), the X-ray spectrum was accumulated from 0.046 to 0.074 days with 2.5 ks exposure.

Both a single ($\beta_X = \beta_O$) and broken PL ($\beta_X - \beta_O = 0.5$, which is fixed) were fit to the SED, and it was found that a cooling break was not needed to fit the data, with a spectral index of $\beta = 1.0 \pm 0.1$ proving to be a good fit. When a cooling break was imposed upon the data, it was always found to occur below the optical data. The paucity of data does not allow us to discriminate between the different extinction curves of the Small Magellanic Cloud (SMC), Large Magellanic Cloud (LMC) and Milky Way (MW) from Pei (1992), so we adopted the SMC template (which has proved to be a suitable fit to the AG SEDs, e.g. Kann et al. 2006). Our best-fitting parameters ($\chi^2/\text{dof} = 29.6/28$) are $A_V = 0.15$ mag (< 0.55 at 90% CL), and an intrinsic column absorption of $N_H = 1.0 \times 10^{21} \text{ cm}^{-2}$ ($< 2.7 \times 10^{21}$ at 90% CL). To convert the rest-frame extinction into equivalent observer-frame extinctions in our SDSS filters we used the SMC extinction template at $z = 0.8$ and the effective wavelengths in Fukugita et al. (1995), finding: $A_{g,\text{obs}} = 0.34$ mag, $A_{r,\text{obs}} = 0.26$ mag and $A_{i,\text{obs}} = 0.20$ mag. We use these values of the rest-frame extinction throughout our analysis of GRB 120729A.

2.4. The Host Galaxy

We used our *griz* observations of the host galaxy, taken at $t - t_0 \approx 189$ d and corrected for foreground extinction, to constrain some of its key physical properties (Fig. 3). Our procedure involves fitting the photometry with stellar population synthesis models from Bruzual & Charlot (2003) with LePHARE (Arnouts et al. 1999). We use a Calzetti dust attenuation law (Calzetti et al. 2000), a Chabrier initial mass function (Chabrier, 2003) and a grid of different star-burst ages with varying e -folding timescales to derive theoretical galaxy spectra which then were compared to our photometry. A more elaborate description of our SED fitting procedure and its caveats is given in Krühler et al. (2011).

The best fitting template is for that of a low-mass, blue, young star-forming galaxy. The best-fitting parameters are: $M_B = -19.3 \pm 0.1$, $\log_{10}(\text{stellar mass}) = 8.3 \pm 0.2 M_\odot$, $\text{SFR} = 6^{+25}_{-4} M_\odot \text{ yr}^{-1}$ and the age of the starburst ≤ 100 Myr. The SFR has a large uncertainty due to the unknown dust attenuation.

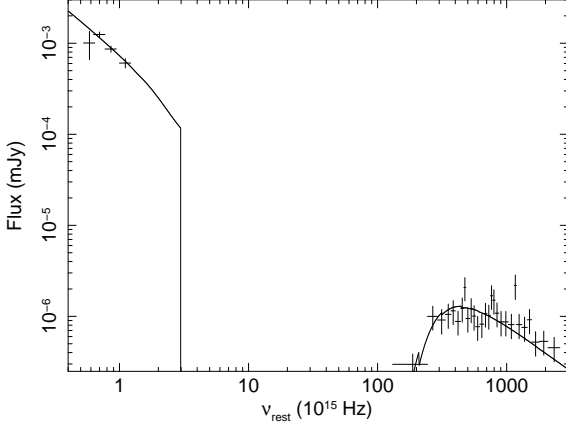


Fig. 2. GRB 120729A: Rest-frame X-ray to optical SED of the AG at $t - t_0 = 0.42$ d. It is found that a single PL provides a good fit to the data, with $\beta = 1.0 \pm 0.1$. Our best-fitting parameters ($\chi^2/dof = 29.6/28$) are $A_V = 0.15$ mag (< 0.55 at 90% CL), and an intrinsic column absorption of $N_H = 1.0 \times 10^{21}$ cm $^{-2}$ ($< 2.7 \times 10^{21}$ at 90% CL).

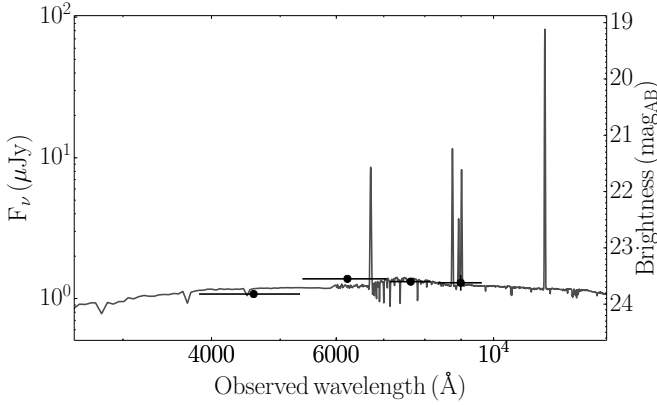


Fig. 3. GRB 120729A: Best-fitting SED of the host galaxy *griz* magnitudes. The best fitting template is for a low-mass ($\log_{10}(\text{mass}) = 8.3 \pm 0.2$ M_\odot), blue, star-forming galaxy (SFR = 6_{-4}^{+25} M_\odot yr $^{-1}$ and the age of the starburst ≤ 100 Myr).

2.5. The Supernova

The dearth of late-time observations limits our analysis of the accompanying SN to GRB 120729A, where only a few detections have been made near peak in *r* and *i*. Indeed the shape of the SN LC is not well constrained, especially given the lack of detections after the peak. Nevertheless, despite this limitation we estimated the brightness of the SN in both filters during our fit. In addition to fitting broken power-laws to LCs we included an additional SN-component. Using the same C-code that was developed and adopted in C13⁵, we generated synthetic, *k*-corrected LCs of a template LC (SN 1998bw) as it would appear if it occurred at $z = 0.80$. Then using Pyxplot⁶ we fit the synthetic LC with a linear spline. This spline is then incorporated into another function (equation 5 in C13) that transforms it by a stretch (*s*) and luminosity (*k*) factor. In events where there are many obser-

vations of the SN bump (e.g. GRB 130427A / SN 2013cq, Xu et al. 2013a) it is possible to constrain both *s* and *k*, however here we fix $s = 1.0$ and allowed only *k* to be a free parameter during the fit. Our best-fitting parameters are: $k_r = 1.29 \pm 0.19$ and $k_i = 0.76 \pm 0.11$. Taking these at face value implies (observer-frame) peak absolute magnitudes of $M_{r,\text{peak}} = -18.96 \pm 0.15$ and $M_{i,\text{peak}} = -19.29 \pm 0.15$, though these are somewhat tentative at best due to the uncertain stretch-factor of the SN, where a larger stretch factor implies a brighter luminosity. With the same reasoning we have not attempted to estimate the time of peak light in each filter.

Using the method presented in C13, we estimated the bolometric properties of the accompanying SN. Given that we have not been able to constrain the shape (i.e. width) of the SN in either filter, there is little merit in estimating its ejecta mass and in turn its kinetic energy. However we can make an estimate of the amount of nickel that was nucleosynthesized during the explosion using “Arnett’s Rule” (Arnett 1982) – i.e. the luminosity at maximum is proportional to the instantaneous energy deposition from the radioactive decay of nickel and cobalt. Making the assumption that the average luminosity factor of the accompanying SN in the optical filters is a suitable proxy for the relative difference in luminosity between this SN and the template, (which was shown in C13 to have an uncertainty of order 10%), and using an average luminosity factor of $\bar{k} = 1.02 \pm 0.26$, and fixing $s = 1.0$, we estimate that in the filter range *UBVRIZH* the accompanying SN has a nickel mass of $M_{\text{Ni}} = 0.42 \pm 0.11$ M_\odot . The quoted error is statistical only, and arises from the uncertainty in the luminosity factor. This nickel mass is close to that estimated for the archetypal GRB-SN 1998bw, where it is estimated 0.4–0.7 M_\odot was nucleosynthesized (Iwamoto et al. 1998; Nakamura et al. 2001; Cano 2013).

3. GRB 130215A

GRB 130215A was detected at 01:31:30 UT on 15-February-2013 by the *Swift*-BAT, and has a $T_{90} = 65.7 \pm 10.8$ s in the 15–350 keV energy range (D’Elia et al. 2013; Barthelmy et al. 2013a). Due to a Moon observing constraint, *Swift* could not slew to the BAT position, thus there are no XRT or UVOT data for this GRB. The burst was also observed by *Fermi*-GBM (Younes & Bhat, 2013) with $T_{90} \approx 140$ s in the 50–300 keV energy range; and by the Sukaku Wide-Band All-sky Monitor (WAM) with $T_{90} \approx 46$ s in the 100–1000 keV energy range (Ishida et al. 2013). Rapid follow-up observations were performed by many ground-based telescopes (Zheng et al. 2013a, 2013b; LaCluyze et al. 2013; Cenko, 2013; Covino et al. 2013; Butler et al. 2013a, 2013b; Gendre et al. 2013; Xu & Zhang, 2013; Hentunen et al. 2013a; Zhao & Bai, 2013; Wren et al. 2013; Kuroda et al. 2013; Knust et al. 2013; Perley, 2013a, 2013b; Singer et al. 2013). The redshift was measured to be $z = 0.597$ (Cucchiara et al. 2013). The AG was clearly detected at 93 GHz at +2.73 hr (Perley & Keating 2013). A spectrum of SN 2013ez was obtained with the GTC (de Ugarte Postigo et al. 2013a, 2013b). An estimate of the isotropic energy release in γ -rays ($1 - 10^4$ keV, rest-frame) is $E_{\text{iso},\gamma} = 3.1_{-1.6}^{+0.9} \times 10^{52}$ erg⁷, and $E_p = 155 \pm 63$ keV (Younes & Bhat 2013). The probability that GRB 130215A arises from a collapsar (Bromberg et al. 2013) based on T_{90} alone is $99.995 \pm 0.002\%$ (BAT) and $99.487 \pm 0.358\%$ (GBM). We use a foreground extinction value of $E(B - V)_{\text{fore}} = 0.162$ mag for GRB 130215A.

⁵ see as well Cano (2014, in prep.)

⁶ <http://pyxplot.org.uk>

⁷ http://butler.lab.asu.edu/swift/bat_spec_table.html

3.1. Data Reduction, Photometry & Spectroscopy

We obtained observations with several ground-based telescopes. ROTSE-III automatically started imaging the field of GRB 130215A 697 s after the initial γ -ray trigger, locating a new, bright (unfiltered=14.2) source at 02:54:00.73 +13:23:43.7 (J2000), with an uncertainty of $< 1''$. Further observations were obtained during the first day with the Nordic Optical Telescope (NOT) and the Gamma-Ray burst Optical/Near-Infrared Detector (GROND; Greiner et al. 2008). Several hours of observations were obtained with the Reionization and Transients Infrared Camera (RATIR⁸) on the 1.5-m Harold Johnson Telescope at the Observatorio Astronómico Nacional on Sierra San Pedro Mártir during the first few hours after the trigger, with additional epochs at $t - t_0 = 2, 3, 4, 8, 11, 17$ d. We also obtained two epochs of spectroscopy and one epoch of optical photometry with the GTC. Early spectroscopy of the AG was performed with OSIRIS, $t - t_0 = 0.79$ d after the GRB using the R1000B grism, which gives a spectral resolution of $\delta\lambda/\lambda \sim 1000$ and a coverage from 3600 to 7500 Å. We obtained an additional spectrum of the accompanying SN 2013ez at $t - t_0 = 25.78$ d, the timing of which was planned to observe the SN at or near maximum light. This observation was performed using the R500R grism, with a spectral resolution of $\delta\lambda/\lambda \sim 600$ and coverage from 4800 to 10000 Å. Each spectrum was reduced using standard techniques with IRAF-based scripts. Late time images ($t - t_0 = 372.8$ d) of the GRB field were obtained with the GTC in filters *gri*, while a late epoch ($t - t_0 = 331.8$ d) was obtained with the 3.5-m CAHA telescope in *J*.

The optical data were calibrated via standard star photometry. On 21-August-2013 GROND obtained images of the GRB field and an SDSS (Abazajian et al. 2009) field located at 03:00:48.0, +19:57:00 (J2000), with the SDSS field taken immediately after the GRB field. Both sets of images were taken under photometric conditions. The calibration was performed using a zeropoint and an airmass correction, and the solution was used to calibrate a set of secondary standards in the field of the GRB. Each dataset was then calibrated to a subset of these stars, depending which ones were in the field of view of each telescope. A summary of our photometry is presented in Table 3.

3.2. The Afterglow

Figure 4 displays our optical photometry, which has been corrected for foreground extinction and then converted into monochromatic fluxes. The LCs were simultaneously fit with a single PL up to $t - t_0 = 1.0$ d (except the Y-band data, which were normalised using the detection at 2.5 d, and is likely overestimated in brightness as this detection appears to be during the plateau phase), with the best-fitting value of the temporal index being $\alpha = 1.25 \pm 0.11$ ($\chi^2/dof = 231.9/141$). After one day the LCs deviate away from the PL-like decline and undergo a plateau phase that lasts up to six days post burst. The LCs then break again before leveling out a further time due to light coming from SN 2013ez.

Due to the lack of XRT data we have not been able to construct an X-ray to optical energy spectrum. Instead we used our contemporaneous optical/NIR data taken with RATIR and GROND over several epochs to get an estimate of the rest-frame extinction. Using the same epochs as those used to construct the bolometric LC in Section 3.3, we fit the empirical extinction curves of the SMC, LMC and MW from Pei (1992), using

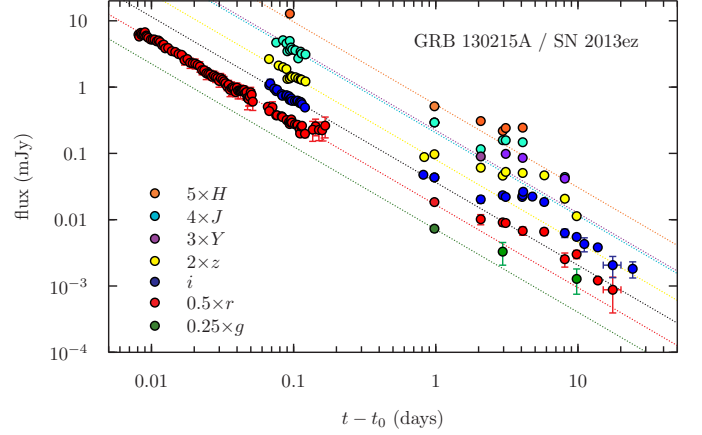


Fig. 4. GRB 130215A: Optical and NIR LCs. The optical data have been corrected for foreground extinction and converted into mJy using the flux zeropoints from Fukugita et al. (1995). All filters have been fit with a single PL up to $t - t_0 = 1.0$ d (except Y which was normalised using the detection at 2.5 d), where $\alpha = 1.25 \pm 0.01$ ($\chi^2/dof = 231.9/141$). A plateau is seen in all filters from $t - t_0 = 1$ –6 d, where each LC deviates away from a single PL-like decline.

a method similar to Kann et al. (2006) and Kann et al. (2010). Each SED is well described by a single PL ($\beta = 0.9 \pm 0.2$ over all epochs), with very little if any curvature, implying there is no need to invoke the presence of rest-frame dust extinction. Each epoch is equally well fit by each dust-extinction template. Indeed some epochs predict a (very small) negative value for the extinction, which is an unphysical conclusion, while the other epochs are consistent with zero rest-frame extinction. Throughout the rest of the analysis of GRB 130215A, we assume $E(B - V)_{\text{rest}} = 0.0$ mag.

3.3. Magnetar Origins?

There are many examples of GRB LCs that show deviations away from a PL-like decay, e.g. GRB 011211 (Jakobsson et al. 2003), GRB 021004 (de Ugarte Postigo et al. 2005), GRB 030429 (Jakobsson et al. 2004), GRB 050502A (Guidorzi et al. 2005), GRB 060526 (Thöne et al. 2010), GRB 090926A (Rau et al. 2010; Cenko et al. 2011) and GRB 100814A (de Pasquale et al. 2013; Nardini et al. 2014). One is also reminded of the peculiar LC of GRB 030329 (Matheson et al. 2003; Lipkin et al. 2004) that displayed very complex behaviour and complicated the decomposition of the SN light from the LC. So while the AG LCs of some GRB-SNe are rather smooth (e.g. GRB 090618, Cano et al. 2011a), others are very complex.

The term “energy injection” is used to explain these peculiar bumps, flares and plateaus, where extra energy is pumped into the FS, causing the AG to become brighter (e.g. Panaitescu et al. 1998; Rees & Mészáros 1998; Kumar & Piran 2000; Sari & Mészáros 2000). Energy injection can arise from different physical sources including Poynting flux emitted by a central engine (e.g. Usov 1992; Dai & Lu 1998), the collision of additional shells of material that collide with the original shells that generated the initial γ -ray burst (Zhang et al. 2006); a reverse shock (RS) created from the collision and pile up of multiple shells with the original shells (e.g. Sari & Mészáros 2000; Kobayashi 2000; Harrison & Kobayashi 2013; Japelj et al. 2014); a two-

⁸ www.ratir.org

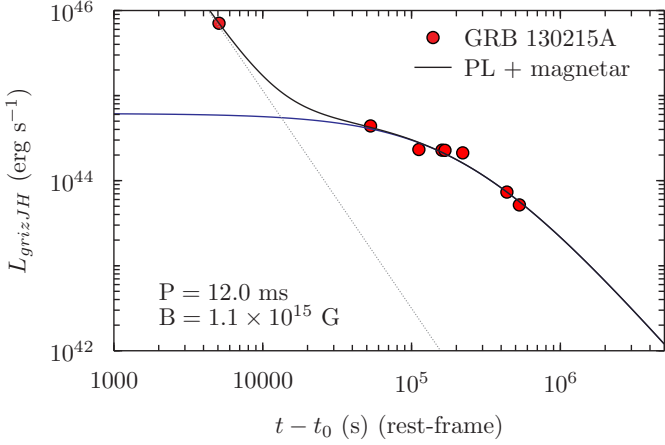


Fig. 5. GRB 130215A: Rest-frame bolometric LC created from our *grizJH* observations. The analytical model from Zhang & Mészáros (2001; see also Rowlinson et al. 2013) has been fit to the LC, which considers energy injection from a millisecond magnetar (plateau and late decline) added to an initial PL-like decline. From the model we find an initial spin period of $P_0 = 12.0$ ms, a magnetic field strength of $B = 1.1 \times 10^{15}$ G, a plateau luminosity of $L_{\text{plat}} = 6.1 \times 10^{44}$ erg s $^{-1}$ and a rest-frame plateau duration of $T_{\text{plat}} = 2.3 \times 10^5$ s (2.7d). Encouragingly, the values of the initial spin period and B -field are *realistic*, and commensurate to those found for long-duration GRBs 060729 and 130427A (see the main text), as well as the sample of short-duration GRBs in Rowlinson et al. (2013).

component jet (Granot et al. 2005) where a rebrightening in the optical bands can arise due to emission from a narrow jet seen off-axis; or a combination of forward and reverse shocks (de Pasquale et al. 2013) including the “thick-shell” scenario, where a combination of the forward and reverse shock (the latter is relativistic) leads to a plateau phase in the observations (Leventis et al. 2014). A more exotic source of energy injection can arise from a quark nova (Staff et al. 2008). During the transition of the newly formed compact object from neutron star \rightarrow quark star \rightarrow black hole, accretion onto the quark star produces a source of extra energy that can be pumped into the ejecta, which can account for the prompt emission as well as flares and plateaus in X-ray LCs. However, injection from an accreting quark star cannot explain plateaus in optical/NIR LCs. One key idea that all these models have in common is that the later the energy injection episode, the more energy is required to create bumps and plateaus of similar magnitude.

Another source of energy injection into the FS can arise from a millisecond magnetar central engine, which deposits Poynting flux dominated dipole radiation into the ejecta (e.g. Zhang & Mészáros 2001; Dall’Osso et al. 2011). The millisecond magnetar model has been considered as a plausible source of energy injection for GRBs, with some notable examples being GRB 000301C (Zhang & Mészáros 2001), GRB 060729 (Xu et al. 2009; Dall’Osso et al. 2011; Lü & Zhang 2014), GRB 120326A (Hou et al. 2014) and GRB 130427A (Bernardini et al. 2014). In these investigations plateau phases in the X-ray LCs are attributed to extra energy arising from a millisecond magnetar, where energy injection refreshes the FS. This is in contrast to the analysis of GRB 070110 (Troja et al. 2007), XRF 100316D (Fan et al. 2011), a sample of long-duration GRBs investigated by Lyons et al. (2010) and the recent study of a sample of short GRBs by Rowlinson et al. (2013), where both authors

attribute the plateaus in the X-ray LCs as flux coming *directly* from the millisecond magnetar.

To our knowledge, to date no attempt has been made to constrain the behaviour of a possible magnetar central engine using a bolometric LC of the AG constructed from optical/NIR observations. Predominantly bolometric X-ray modelling has been the status quo, though an estimate of the optical contribution (R -band) was made for GRB 130427A by Bernardini et al. (2014) and R -band data of GRB 120326A (Hou et al. 2014). In this work we are able to fully exploit the wide filter coverage and simultaneous observations obtained by both GROND and RATIR to create a bolometric LC in the filter range *grizJH* (observer frame) with the aim of determining whether energy injection from a magnetar central engine provides a plausible explanation for the plateaus seen in the optical/NIR LCs. We used data from a total of eight epochs (ranging from $t - t_0 = 0.1$ –9.8 d in the observer frame); i.e. before the period where the SN starts to dominate the LCs). We then followed a standard method to construct our bolometric LC (e.g. Cano et al. 2014): (1) correct all magnitudes for foreground and rest-frame extinction, (2) convert magnitudes into monochromatic fluxes using flux zeropoints in Fukugita et al. (1995). For epochs where there are no contemporaneous observations, we linearly interpolated the flux LCs and SEDs to estimate the missing flux. Then, for each epoch of multi-band observations, and using the effective wavelengths from Fukugita et al. (1995) we: (3) interpolate (linearly) between each datapoint, then (4) integrate the SED over frequency, assuming zero flux at the integration limits, and finally (5) correct for filter overlap. The linear interpolation and integration were performed using a program written in Pyxplot. The resultant LC is shown in Figure 5.

We make similar assumptions as Rowlinson et al. (2013), namely that the magnetar mass is $1.4 M_{\odot}$ and its radius is 10^6 cm, which allows us to reduce the number of free parameters in the fit. The final fit is a combination of an initial PL added to the magnetar model:

$$L_{\text{magnetar}}(t) = L_0 \left(1 + \frac{t}{T_0} \right)^{-2} + \Lambda t^{-\alpha} \quad (1)$$

where L_0 is the plateau luminosity, T_0 is the plateau duration, and Λ is the normalisation constant for the PL. The values of L_0 and T_0 can be related back to equations 6 and 8 in Zhang & Mészáros (2001; see as well Rowlinson et al. 2013) to estimate the initial spin period and magnetic field strength of the magnetar. Fitting this model to our rest-frame bolometric LC, we find an initial spin period of $P_0 = 12.0$ ms, a magnetic field strength of $B = 1.1 \times 10^{15}$ G, a plateau luminosity of $L_0 = 6.1 \times 10^{44}$ erg s $^{-1}$, a rest-frame plateau duration of $T_{\text{plat}} = 2.3 \times 10^5$ s, and $\alpha = 2.6 \pm 0.7$.

Encouragingly, the values of the initial spin and magnetic field are *realistic*, and are found to be comparable to those found for other GRBs with associated SNe: (1) GRB 060729: $P = 1.5$ ms and $B = 0.27 \times 10^{15}$ G (Xu et al. 2009); $P = 2.0$ ms and $B = 3.2 \times 10^{15}$ G (Dall’Osso et al. 2011); and $P = 1.5$ ms and $B = 0.25 \times 10^{15}$ G (Lu & Zhang 2014); (2) GRB 130427A: $P \sim 20$ ms and $B \sim 10^{16}$ G (Bernardini et al. 2014). The spin period determined from observations of GRB 130215A falls within the estimates for GRBs 060729 and 130427A, while the magnetic field strengths vary by two orders of magnitude for these three events. Moreover, these values are fully consistent with the values determined for a sample of long GRBs by Lyons et al. (2010; their Table 2), a sample of short GRBs by Rowlinson et al. (2013; their Table 3), and for short GRB 130603B (de Ugarte

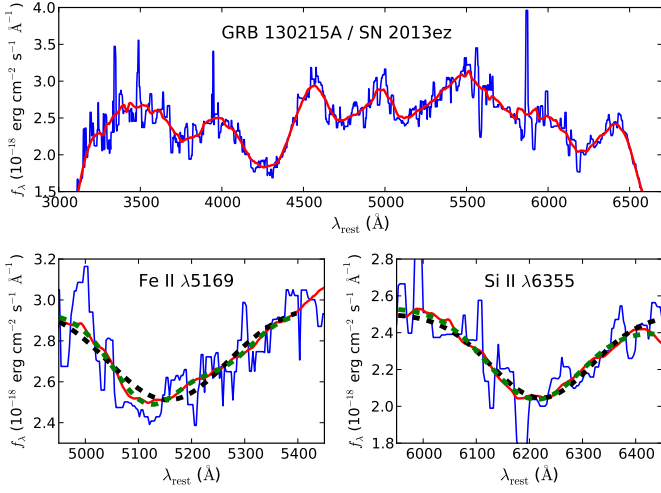


Fig. 6. GRB 130215A / SN 2013ez: SN+host+AG (blue) rest-frame spectrum (λ in Å), from $t - t_0 = 25.8$ d (16.1 d rest-frame). *Top:* The spectrum in blue has been median filtered (width of 20 Å), while the red spectrum has been smoothed using a Kaiser window ($M = 40, \beta = 0$). Clear undulations are seen in the spectrum that are reminiscent of other SNe Ic. *Bottom left:* Close-up of the absorption feature near 5100 Å, which is thought to be blueshifted Fe II $\lambda 5169$. Both a single and double Gaussian have been fit to the smoothed spectrum. The latter provides a better fit ($\lambda_1 = 5100.59 \pm 24.45$ Å, $v_{\text{Fe}} = -4000 \pm 1496$ km s $^{-1}$; $\lambda_2 = 5305.72 \pm 202.17$ Å, which is too red to be Fe II $\lambda 5169$). *Bottom right:* Close-up of the absorption feature near 6200 Å, which is thought to be blueshifted Si II $\lambda 6355$. A single Gaussian was fit to the spectrum, yielding $\lambda_1 = 6221.75 \pm 16.77$ Å, $v_{\text{Fe}} = -6354 \pm 808$ km s $^{-1}$. A double Gaussian was also fit to the smoothed spectrum, however an improved fit was not obtained. The blueshifted velocities of both lines are comparable within their respective error bars, and show that SN 2013ez may be more appropriately classified as a type Ic SN rather than a type Ic-BL.

Postigo et al. 2014). Further discussion on the plausibility of energy injection arising from a millisecond magnetar is presented in Section 5.

3.4. The Host Galaxy

We re-observed the field of GRB 130215A on 22-February-2014 with the GTC telescope in filters *gri*. The host galaxy is not visible in any of our co-added images. We derive 3- σ upper limits for an isolated point source in our images of: $g > 26.2$, $r > 26.1$, $i > 25.1$, which are not corrected for foreground extinction. We also obtained a late-time J-band image with the 3.5-m CAHA telescope on 12-January-2014, where again no object is detected at the position of the GRB. We derive an upper limit of $J > 23.2$.

At $z = 0.597$, and a distance modulus of $\mu = 42.72$, these upper limits imply observer-frame, absolute magnitude limits of the host galaxy of $M_g > -16.5$, $M_r > -16.6$, $M_i > -17.6$ and $M_J > -19.5$.

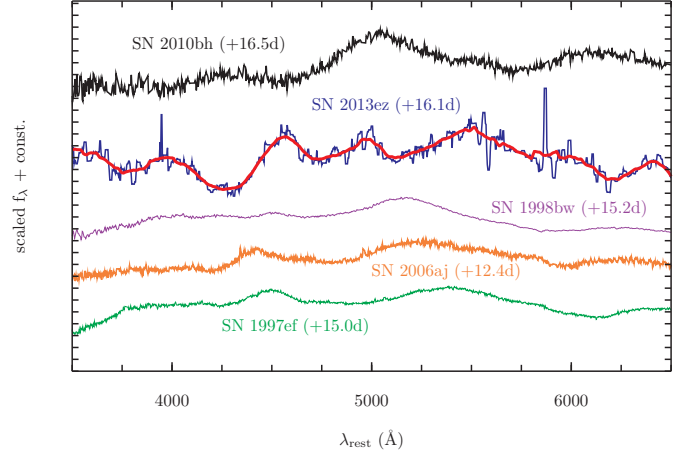


Fig. 7. GRB 130215A / SN 2013ez: For comparison, we plot the spectra of several GRB-SNe (SN 1998bw, purple line; SN 2006aj, orange line; SN 2010bh, black line) and one SN IcBL (SN 1997ef, green line) that was not associated with a GRB, each at a similar post-explosion date as SN 2013ez. All times are rest-frame and the spectra have been scaled and shifted for clarity. It is seen that the absorption features in the spectrum of 2013ez are not as broad as those of the comparison SNe.

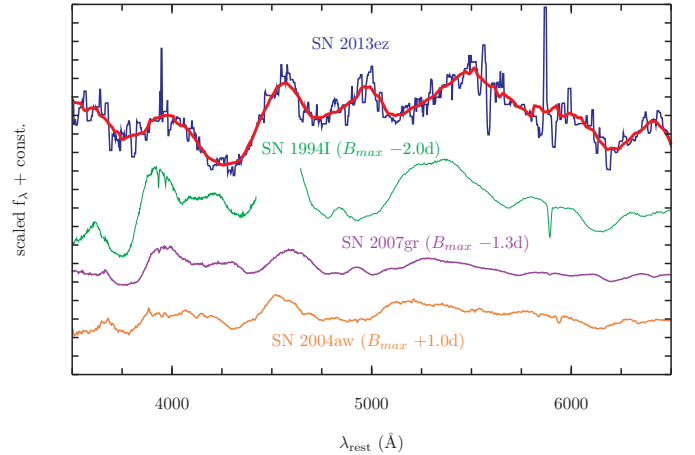


Fig. 8. GRB 130215A / SN 2013ez: Plotted for comparison are the rest-frame spectra of several Ic SNe: 1994I (green), SN 2004aw (orange) and SN 2007gr (purple). All times are from the time of maximum *B*-band light for each SN. The narrow features of SN 2013ez are more reminiscent of those of SNe Ic.

3.5. The Supernova

3.5.1. Spectroscopy of SN 2013ez

The rest-frame (wavelength) spectrum of SN 2013ez at $t - t_0 = 25.8$ d (16.1 d rest-frame) is displayed in Fig. 6. The spectrum in blue is the result of applying a moving median filter (width=20 Å) to the original spectrum. In order to more clearly distinguish the key absorption features, we smoothed the spectrum using a Kaiser window as a smoothing kernel, where the size of the window (M) and shape of the window (β) were varied. The smoothing procedure was performed using a Python program that uti-

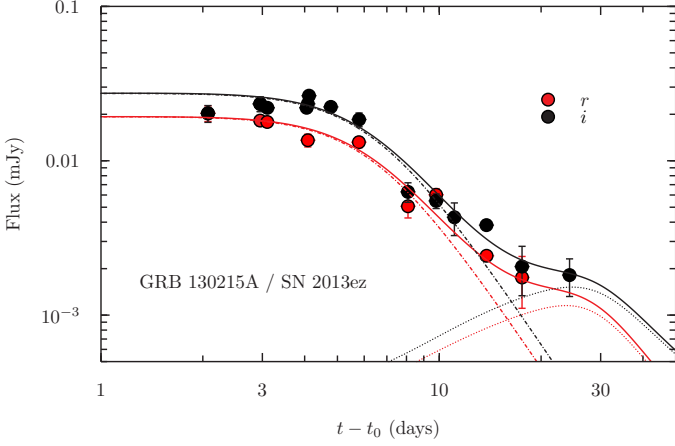


Fig. 9. GRB 130215A / SN 2013ez: Optical LCs in r (red) and i (black). The AG (dot-dashed) and SN (dotted) components are shown in the same colour as their corresponding filter, while the solid lines are the sum of both components. Data at times $> +2$ d have been fit with a broken PL consisting of a plateau phase and a break to a steeper decay phase. We assume that the time at which the LC breaks and the decay rate after the break are the same in both filters. The best-fitting values (for magnitudes that are not host-subtracted, see the main text) are: $\alpha_2 = 3.28 \pm 0.25$ and $T_B = 6.39 \pm 0.35$ d. Due to the (1) lack of host detection in our deep GTC images, and (2) lack of datapoints at times when the SN is the dominant source of flux we have not been able to precisely constrain the SN’s properties. When we consider the two extremes of the host brightness (see main text), we constrain the luminosity factor of SN 2013ez to be $0.6 \leq k \leq 0.75$.

lizes numpy to convolve the spectrum by the Kaiser window. In Fig. 6 values of $M = 40$ and $\beta = 0$ are used.

Two absorption features are seen near 5100 \AA and 6200 \AA , which are thought to be blueshifted Fe II $\lambda 5169$ and Si II $\lambda 6355$ respectively. We fit a single and double Gaussian to both features in the smoothed spectrum in order to determine their central wavelengths and blueshifted velocities (see bottom insets in Fig. 6). In order to get an estimate of the error of each wavelength (and velocity) measurement, we performed a Monte Carlo (MC) simulation to generate 10,000 spectra from the original spectrum and its error spectrum. At each wavelength in the original spectrum, we derived a random number from a Gaussian distribution that is centered at each wavelength in the spectrum, and whose standard deviation is equal to the value of the error spectrum at that wavelength. We also allowed M and β to be random numbers between two pre-determined values. Then, in the MC simulation we fit both a single (SG) and double Gaussian (DG) to the Fe II and Si II features, calculating their central wavelengths and blueshifted velocities. Presented below are the average values of the 10,000 computed wavelengths and velocities, where the $1\text{-}\sigma$ errors have been calculated from their respective standard deviations. The errors are a combination of those arising from the error spectrum and the fitting procedure. The systematic error in λ and the blueshifted velocities arising from fitting the absorption features in a single spectrum 10,000 times are $\approx 2 \text{ \AA}$, and $\approx 100 \text{ km s}^{-1}$ respectively. The errors below are clearly dominated by the error spectrum.

The results of our MC simulation for Fe II $\lambda 5169$ are: (SG): $\lambda = 5125.27 \pm 24.90 \text{ \AA}$; $v = -2552 \pm 1464 \text{ km s}^{-1}$, (DG): $\lambda_1 = 5100.59 \pm 25.45 \text{ \AA}$; $v_1 = -4000 \pm 1496 \text{ km s}^{-1}$; $\lambda_2 = 5305.72 \pm$

202.17 \AA . λ_2 is clearly too red to be Fe II $\lambda 5169$, and so we have not computed its blueshifted velocity.

Similarly, for Si II $\lambda 6355$: (SG): $\lambda = 6221.75 \pm 16.77 \text{ \AA}$; $v = -6357 \pm 808 \text{ km s}^{-1}$, (DG): $\lambda_1 = 6172.44 \pm 569.46 \text{ \AA}$; $v_1 = -8284 \pm 6219 \text{ km s}^{-1}$; $\lambda_2 = 6156.27 \pm 7406.63 \text{ \AA}$; $v_2 = -1643 \pm 11,544 \text{ km s}^{-1}$. Clearly the double Gaussian does not improve the fit, and the central wavelengths of both are approximately equal.

In summary, the Fe II $\lambda 5169$ feature is better described by a DG, where a blueshifted velocity of $v \approx -4000 \pm 1500 \text{ km s}^{-1}$ is found. The Si II $\lambda 6355$ is better described by a SG, where a blueshifted velocity of $v \approx -6350 \pm 800 \text{ km s}^{-1}$ is found. The velocities, of quite modest values, agree within their respective errorbars.

To put these velocities into context, Schulze et al. (2014) investigated the velocity of Fe II $\lambda 5169$ for a sample of seven GRB-SNe. At $t - t_0 = 16$ d (rest-frame), the velocity of this line ranges from $\sim -14,000 \text{ km s}^{-1}$ for SN 2006aj to $\sim -36,000 \text{ km s}^{-1}$ for SN 2010bh. The velocity of this line in the spectrum of SN 2013ez is $\sim 8000\text{--}10,000 \text{ km s}^{-1}$ slower than that of SN 2006aj, which itself represents the slowest velocity among all the GRB- and XRF-SNe in their sample. Based on this argument it may be more appropriate to classify SN 2013ez as type Ic rather than Ic-BL.

Plotted for comparison⁹ in Fig. 7 are the spectra of SN 1998bw (Patat et al. 2001) that was associated with GRB 980425, SN 2006aj (Modjaz et al. 2006) that was associated with XRF 060218, SN 2010bh (Bufano et al. 2012) that was associated with XRF 100316D, and SN 1997ef (Garnavich et al. 1997; Hu et al. 1997) which is a SN IcBL not associated with a GRB. Each spectrum is taken near a similar rest-frame time to that of SN 2013ez: SN 1998bw ($t - t_0 = 15.2$ d; Patat et al. 2001), SN 2006aj ($t - t_0 = 12.4$ d; Modjaz et al. 2006), SN 2010bh ($t - t_0 = 16.5$ d; Bufano et al. 2011), and SN 1997ef ($t - t_0 \approx 15$ d; Mazzali et al. 2000). The majority of the telluric and host galactic lines have been removed, and the spectra have been shifted and scaled to provide a clear comparison. It is seen that the absorption features in the spectrum of SN 2013ez are not as broad as those of the comparison SNe.

Conversely, plotted in Fig. 8 are three Ic SNe not associated with a GRB near maximum B -band light: SN 1994I ($B_{\text{max}} - 2$ d; Filippenko et al. 1995), SN 2004aw ($B_{\text{max}} + 1$ d; Taubenburger et al. 2006) and SN 2007gr ($B_{\text{max}} - 1.3$ d; Valenti et al. 2008). The narrow lines in the comparison SNe Ic are redolent of those of SN 2013ez.

3.5.2. Photometry of SN 2013ez

While the spectroscopic identification of SN 2013ez is unambiguous, the plateau in the optical/NIR complicates our ambition of decomposing the LCs in order to isolate the photometric SN contribution. The situation is also perplexed by the paltriness of photometric observations of SN 2013ez near peak. Nevertheless we decomposed the optical LCs to estimate the brightness of SN 2013ez in filters r and i (Fig. 9). We fit a broken PL to the LCs, and imposed a plateau phase ($\alpha_1 = -0.01$), which breaks at some time (T_B) to a steeper decay phase (α_2). We assume that the time the LC breaks and the rate of decay after the break are the same in both filters, and these two parameters are allowed to vary during the fit.

⁹ The majority of the comparison spectra were downloaded from the WiseREP SN spectrum database (Yaron & Gal-Yam, 2012), with the rest coming (with our gratitude) from their respective references.

The decomposition is further complicated by the fact that we have not detected the host in our deep GTC images (see section 3.4), so we must consider two scenarios: (1) magnitudes that are not host-subtracted, and (2) take the host brightness equal to the limits obtained from the GTC images (and correct for foreground extinction). These two scenarios can be considered to be the two extremes to the SN brightness, for certainly the host will be contributing *some* flux, but no more than the upper limits of the GTC images.

In scenario (1) we find best-fitting AG parameters of $\alpha_2 = 3.28 \pm 0.25$ and $T_B = 6.39 \pm 0.35$ d, while in scenario (2) we find $\alpha_2 = 3.44 \pm 0.28$ and $T_B = 6.41 \pm 0.34$ d. Unsurprisingly the time the LC breaks is essentially the same in both scenarios, while the LCs decay faster when we remove a host contribution. We also note that the break-time is later than that found in the magnetar model ($T_B = 4.3$ d in observer frame). As for GRB 120729A, we fix the stretch factor to $s = 1.0$ due to the lack of datapoints. In scenario (1) we find $k \approx 0.75$, while in scenario (2) we find $k \approx 0.6$. These two values can be considered the upper and lower limits to the brightness of SN 2013ez in these filters. Taking these values at face value implies peak brightnesses of $M_r = -18.7$ to -19.0 , and $M_i = -19.0$ to -19.3 . Again, there is little merit in estimating the peak times due to the unknown stretch values. Finally, using the method in C13, we estimate a nickel mass in the range $0.25 \leq M_{\text{Ni}} \leq 0.30 M_{\odot}$.

4. GRB 130831A

GRB 130831A was detected at 13:04:16 UT on 31-August-2013 by the *Swift* Burst Alert Telescope (BAT), and has a $T_{90} = 32.5 \pm 2.5$ s in the 15–350 keV energy range (Hagen et al. 2013; Barthelmy et al. 2013b). It was also detected by Konus-Wind, Golenetskii et al. (2013) estimate an isotropic energy release in γ -rays of $E_{\text{iso},\gamma} = 4.6 \pm 0.2 \times 10^{51}$ erg in the 20 keV–10 MeV range, and $E_p = 67 \pm 4$ keV (Golenetskii et al. 2013). The probability that GRB 130831A arises from a collapsar (Bromberg et al. 2013) based on T_{90} alone is $99.969 \pm 0.006\%$ (BAT).

Rapid follow-up of GRB 130831A was performed by several ground-based telescopes (Guidorzi & Melandri 2013; Xu et al. 2013b; Yoshii et al. 2013; Xin et al. 2013; Trotter et al. 2013; Leonini et al. 2013; Masi & Nocentini 2013; Izzo & D’Avino 2013; Hentunen et al. 2013b; Sonbas et al. 2013; Butler et al. 2013c; Chester & Hagen 2013; Volnova et al. 2013a, 2013b, 2013c, 2013d; Pozanenko et al. 2013 and Khorunzhev et al. 2013). The AG was not detected at radio (Laskar et al. 2013) or sub-mm wavelengths (Zauderer et al. 2013; Smith et al. 2013). A redshift of $z = 0.479$ was measured with Gemini-North (Cucchiara & Perley 2013). A spectrum of the associated supernova, SN 2013fu, was obtained with the VLT by Klose et al. (2013). We use a foreground extinction value of $E(B - V)_{\text{fore}} = 0.046$ mag for GRB 130831A.

4.1. Data Reduction & Photometry

We obtained optical observations with several ground-based telescopes. The 0.65-m SANTEL-650 and 0.5-m VT-50 telescopes of the UAFO/ISON-Ussuriysk started imaging (unfiltered) the GRB field just over 10 minutes after the initial trigger, obtaining nearly consecutive images for six straight hours. Additional follow-up observations were obtained with the Gissar observatory 0.7-m telescope, the 0.4-m SANTEL-400AN telescope (ASC/ISON-Kislovodsk observatory), the 0.7-m AZT-8 telescope operated by the Institute of Astronomy,

Kharkiv National University and the 1.5-m AZT-33IK telescope at Mondy observatory, Shajn 2.6-m telescope of Crimean Astrophysical observatory and the 1.5-m AZT-22 telescope at Maidanak observatory. Data obtained at times $t - t_0 < 2.0$ d with aforementioned Russian telescopes are presented in de Pasquale et al. (2014, in prep), while everything at this time and later are presented in this paper. We obtained several epochs of photometry with the 2.5-m NOT, three epochs with the 4.2-m William Hershel telescope (WHT), and four epochs with the 2.0-m LT. We also obtained a single late-time epoch with the Gemini Multi-Object Spectrograph (GMOS; Hook et al. 2004) mounted at Gemini-South as a part of the program GS-2013B-Q-69. The data were reduced in a standard fashion with the Gemini IRAF software package for GMOS (v1.12).

The optical data were calibrated using SDSS stars in the GRB field with a zeropoint between the instrumental and catalog magnitudes. Observations obtained with each Russian telescope in Johnson/Cousins filters BVR_cI_c were calibrated by converting the SDSS (AB) magnitudes of local standards into Johnson/Cousins (Vega) using transformation equations in Lupton (2005). The late-time r observations taken with the NOT, WHT and Gemini were converted into R_c using transformation equations from Jordi et al. (2006), which require a colour term ($r - i$) in the calculations. In an identical procedure as for GRB 120729A (see section 2.1) we interpolated the i -band LC to the times of the r -band LC, extracting the i magnitude. A summary of our photometry is presented in Table 3.

4.2. The Afterglow

Our R_ciz optical data are displayed in Fig. 11. All magnitudes are corrected for foreground and rest-frame extinction. We host-subtracted the optical data in R_c and i using the host detections (see Section 4.4) in the same filters, and then converting all magnitudes into monochromatic fluxes and then mathematically subtracting the host flux from the earlier epochs. The z data have not been host-subtracted due to lack of observations of the host in z at late times. All LCs are well described by a single PL, where we assume that the LCs decay at the same rate in all filters, where $\alpha = 1.63 \pm 0.02$. We note the presence of a bump or short plateau phase in the R -band LC between $t - t_0 \sim 3$ –5 d, however this short phase does not appear to affect our analysis of the decay rate and subsequent optical properties of the SN in R_c and i .

4.3. The Spectral Energy Distribution

In an identical analysis as Section 2.3 we constructed rest-frame X-ray to optical SEDs in order to get an estimate of the rest-frame extinction (Fig. 10). We fit two epochs of data at $t - t_0 = 0.39$ d and 1.12 d (observer frame), which coincide with optical data obtained with the NOT and MAO respectively. The X-ray spectra were derived from 0.11–0.61 d with a total exposure of 8 ks and rescaled at both epochs of optical data. The optical data have been corrected for foreground extinction.

As before both single and broken PLs were fit to the SEDs, and we find that for both epochs a single PL fits the data well. Our results for both epochs are: (1) $t - t_0 = 0.39$ d ($\chi^2/\text{dof} = 49.2/45$): $\beta = 0.85 \pm 0.01$, $A_V < 0.1$ mag (90% CL), and $N_H = 4.2 \pm 0.8 \times 10^{20} \text{ cm}^{-2}$; (2) $t - t_0 = 1.12$ d ($\chi^2/\text{dof} = 42.0/45$): $\beta = 0.75 \pm 0.06$, $A_V = 0.21^{+0.28}_{-0.21}$ mag, and an intrinsic column density of $N_H = 3.5 \pm 1.0 \times 10^{20} \text{ cm}^{-2}$. We thus conclude that the extinction local to GRB 130831A is consistent with being zero, and for our analysis we use the value of $E(B - V)_{\text{rest}} = 0.0$ mag.

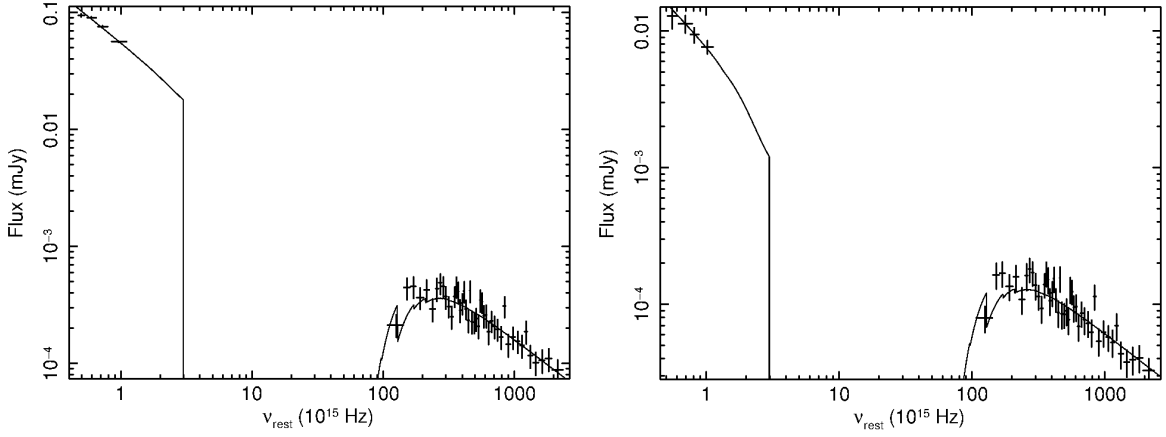


Fig. 10. GRB 130831A: X-ray to optical SED of the AG at $t - t_0 = 0.39$ and 1.12 d (observer frame). It is found that a single PL provides a good fit to both epochs of data ($\beta = 0.85 \pm 0.01$ and 0.75 ± 0.06 respectively). For the first epoch (left) we find $A_V < 0.1$ mag (90% CL), while in the second epoch (right) we find $A_V = 0.21^{+0.28}_{-0.21}$ mag. The results of the SED fitting indicate that the rest-frame extinction is consistent with being $A_V = 0.0$ mag. Additionally we find the intrinsic column density to be $N_H \approx 3 - 4 \times 10^{20} \text{ cm}^{-2}$ for both epochs.

4.4. The Host Galaxy

We observed the field of GRB 130831A at late times with the LT (i) and NOT (r), and find an extended object at the GRB position in both images, which we attribute as the host galaxy. In the i -band LT image taken on 05-Jan-2014 at $t - t_0 = 127.3$ d (+86.1 d in rest-frame), we measure $i = 24.23 \pm 0.10$. In our r -band NOT image taken on 03-Feb-2014 at $t - t_0 = 156.3$ d (+105.7 d in rest-frame), we measure $r = 24.06 \pm 0.09$. These magnitudes are not corrected for foreground extinction. In terms of absolute magnitude, we find (observer-frame) $M_r = -18.06 \pm 0.09$ and $M_i = -17.89 \pm 0.10$. The colour $r - i = -0.17$ suggests that the host galaxy is rather blue. We note that we did not attempt to fit galaxy SEDs due to the sparseness of host observations.

4.5. The Supernova

A clear SN bump in filters riz , which was initially suggested as the signature of SN 2013fu (Pozanenko et al. 2013b), is seen in Fig. 11. The SN bump is particularly pronounced in the well-sampled R -band LC, and is also seen in i , and to a lesser extent in z , where a flattening of the LC is seen, which can be attributed to flux coming from SN 2013fu. When fitting the optical data in Section 4.2 we simultaneously fit a SN-component to determine the stretch (s) and luminosity (k) factors in each filter. Our best-fitting parameters are: $k_R = 0.65 \pm 0.03$ and $s_R = 0.84 \pm 0.03$; $k_i = 1.04 \pm 0.05$ and $s_i = 0.81 \pm 0.03$; $k_z = 1.02 \pm 0.19$ and $s_z = s_i = 0.81$ (fixed). Due to the lack of observations in z at late times we fix the value of s to be the same as in i . Moreover, as the z data are not host-subtracted, the luminosity factor is an upper limit to the maximum brightness of SN 2013fu in this filter.

We determined the peak absolute magnitude, and time of peak light, of SN 2013fu in each filter. In R_c : $M_R = -18.89 \pm 0.05$ and $t_p = 18.60 \pm 0.67$ d (12.58 ± 0.45 d in rest frame); i : $M_i = -19.56 \pm 0.05$ and $t_p = 18.53 \pm 0.069$ d (12.53 ± 0.47 d in rest frame); z : $M_z = -19.47 \pm 0.19$ and $t_p \approx 19.1$ d (≈ 12.9 d in rest frame). The peak time in z is tentative however, given that we have not been able to directly determine the stretch factor directly from our observations. At $z = 0.479$, observer-frame i ($\lambda_{\text{eff}} = 7706 \text{ \AA}$) is roughly rest-frame V ($\lambda_{\text{eff}} = 5505 \text{ \AA}$): $7706/1.4791 = 5210 \text{ \AA}$. Making a k-correction using the formula

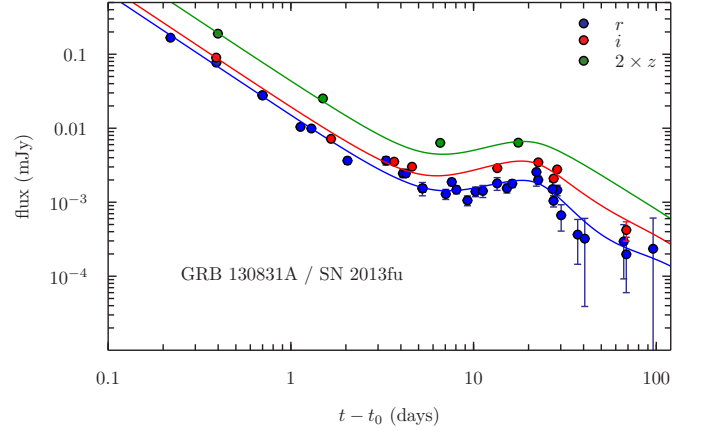


Fig. 11. GRB 130831A: Optical ($R_c iz$) light curves. The solid lines in each filter are the sum of the AG and SN components. The optical data have been corrected for foreground extinction, and the R_c and i are host-subtracted (see the text). The z data have not been host-subtracted due to lack of observations of the host in z at late times. All LCs are well fit with a single PL, and we assume that the LCs decay at the same rate in all filters, where $\alpha = 1.63 \pm 0.02$. A clear SN bump is seen in R_c and i , and a flattening of the LC is seen in z , which can be attributed to flux coming from SN 2013fu. In each filter we simultaneously fit a SN-component to determine the stretch (s) and luminosity (k) factors in each filter. Due to the lack of observations in z at late times we fix the value of stretch factor to be the same in i and z (i.e. $s_i \equiv s_z$). Our best-fitting parameters are: $k_R = 0.65 \pm 0.03$ and $s_R = 0.84 \pm 0.03$; $k_i = 1.04 \pm 0.05$ and $s_i = 0.81 \pm 0.03$; $k_z = 1.02 \pm 0.19$ and $s_z = s_i = 0.81$ (fixed). As the z data are not host-subtracted, the luminosity factor is an upper limit to the maximum brightness of SN 2013fu in this filter.

lation of Hogg et al. (2002), and a spectrum of SN 1998bw as a template, we find a k-correction from observer-frame i to rest-frame V of $k_{i \rightarrow V} \approx 0.25$ mag. This implies a rest-frame, peak magnitude of $M_V \approx -19.31$. This value is consistent with the average peak V-band magnitude found for a sample of k-corrected

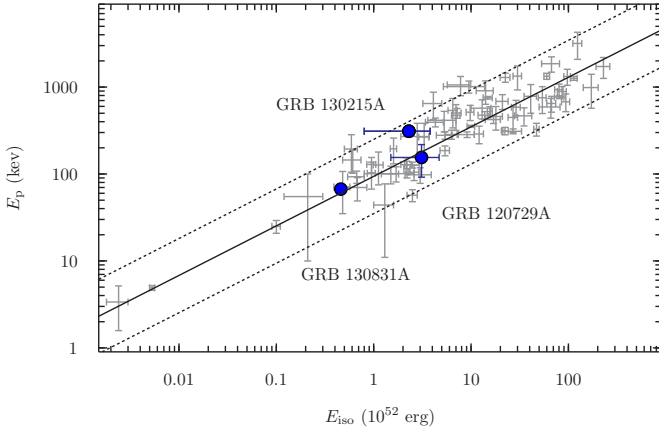


Fig. 12. Our trio of GRBs on the $E_p - E_{\text{iso},\gamma}$ plane. The sample of GRBs presented in Amati et al. (2008) are shown in grey, along with their best fit to a single powerlaw ($\alpha = 0.57$) and the $2\text{-}\sigma$ uncertainty in their fit. GRBs 120729A and 130831A lie perfectly on the Amati relation, while GRB 130215A is within the $2\text{-}\sigma$ confidence interval.

LCs of GRB-SNe analysed by Richardson (2009), who found $M_{V,\text{peak}} = -19.2 \pm 0.2$ (standard deviation of $\sigma = 0.7$ mag).

It is seen that in the well-sampled R -band LC that the SN appears to decrease in brightness faster than the k -corrected LC of SN 1998bw. There is also a hint of this in the I -band LC, though we cannot draw many conclusions based on a single datapoint at late times. When we calculate Δm_{15} in R_c (where Δm_{15} is the amount the LC fades from peak light to 15 days later; which is calculated for rest-frame times) for SN 2013fu and SN 1998bw, (where the latter is transformed by $k_R = 0.65 \pm 0.03$ and $s_R = 0.82 \pm 0.03$), we find $\Delta m_{15} \approx +1.99$ and $\Delta m_{15} \approx +1.45$ respectively. This clearly shows that 2013fu evolves faster than the archetype GRB-SN 1998bw. Therefore, in this case the shape of the template SN does not provide the best description for the temporal evolution of SN 2013fu. This type of behaviour has been seen for other GRB-SNe, such as SN 2010bh associated with XRF 100316D (Cano et al. 2011b), SN 2006aj associated with XRF 060218 (Ferrero et al. 2006), as well as local SNe Ibc presented in C13.

Using the model in C13 we estimated the nickel mass, ejecta mass and kinetic energy of SN 2013fu. Without knowledge of the peak photospheric velocity of SN 2013fu, we used the average peak photospheric velocity determined by C13 for a sample of GRB-SNe: $v_{ph} \approx 20 \pm 2.5 \times 10^3 \text{ km s}^{-1}$. To estimate the nickel mass we computed the average luminosity factor from the r and i filters (neglecting the z observation as it is not host subtracted and is therefore an overestimate of the SN's brightness), $\bar{k} = 0.85 \pm 0.20$. We estimated the ejecta velocity using the peak photospheric velocity and an average of the stretch factor in R and i , $\bar{s} = 0.83 \pm 0.02$. We find bolometric properties of: $M_{\text{Ni}} = 0.30 \pm 0.07 M_{\odot}$, $M_{\text{ej}} = 4.71^{+0.79}_{-0.59} M_{\odot}$ and $E_K = 1.87^{+0.90}_{-0.62} \times 10^{52} \text{ ergs}$. The uncertainties in the ejecta mass and kinetic energy arise from the uncertainties in the stretch and luminosity factors as well as the spread of peak ejecta velocities around the mean value in C13.

5. Discussion & Conclusions

5.1. The Supernovae

We presented optical/NIR photometry for three GRB-SNe. We derived the luminosity factor (k) for each SN, and the stretch factor (s) of SN 2013fu relative to a template supernova (SN 1998bw), which has been redshifted/ k -corrected to that of each GRB-SN considered here. We also estimated the peak, observer-frame magnitude of each SN in every available filter, as well as the time of peak light for SN 2013fu.

We also presented a spectrum of SN 2013ez, which was associated with GRB 130215A. Absorption features near 5100 \AA and 6200 \AA are seen, which are thought to be blueshifted Fe II $\lambda 5169$ and Si II $\lambda 6355$ respectively. Using a Monte Carlo (MC) simulation we fit single and double Gaussians to these features, finding blueshifted velocities of $v_{Fe} \approx -4000 \pm 1500 \text{ km s}^{-1}$ and $v_{Si} \approx -6350 \pm 800 \text{ km s}^{-1}$. These velocities agree within their respective errorbars.

The spectrum was taken at rest-frame $t - t_0 = 16.1 \text{ d}$, which despite the unknown time of peak light, is likely to be no more than a few days from peak light if it evolved similarly to SN 1998bw, or \sim a week from peak light if it evolved at a faster rate such as seen for SN 2006aj and SN 2010bh. The blueshifted velocity seen in SN 2013ez is $\sim 8000\text{--}10,000 \text{ km s}^{-1}$ less than that of the slowest velocity of a sample of seven GRB/XRF-SNe investigated by Schulze et al. (2014) at the same rest-frame time. The slowest blueshifted velocity of Fe II $\lambda 5169$ at $t - t_0 = 16 \text{ d}$ (rest-frame) in their sample is $\sim 14,000 \text{ km s}^{-1}$ for SN 2006aj. Thus the narrow absorption features along with the slower blueshifted velocity of SN 2013ez makes it more appropriate to consider SN 2013ez as a type Ic rather than a type IbL, making it the first GRB-SN to have this classification.

The only other GRB-SN to have a possible classification of Ic was SN 2001lt, which was associated with GRB 021211 (Della Valle et al. 2003). For this SN, blueshifted Ca II was seen at $\sim 14,400 \text{ km s}^{-1}$ at $t - t_0 \approx 13.4 \text{ d}$ (rest-frame), which, while being a lower value than seen in the majority of the other GRB-SNe, is comparable to SN 2006aj at a similar rest-frame epoch. Moreover, the average peak photospheric velocity of a sample of seven SNe IbL investigated by C13 is $\sim 15,000 \text{ km s}^{-1}$ ($\sigma = 4000 \text{ km s}^{-1}$), which is commensurate to that of SN 2001lt, making it comfortably classified as a type Ic-BL rather than Ic.

When analysing the optical properties of SN 2013fu, associated with GRB 130831A, we found that it is brighter in the redder filters: $k_R = 0.65 \pm 0.03$, $k_i = 1.04 \pm 0.05$ and $k_z = 1.02 \pm 0.19$.¹⁰ If we take our result at face value, the red colour of 2013fu suggests that there is a suppression of flux in observer-frame R -band ($\approx B$ -band in rest-frame, e.g. $\lambda_{\text{eff}} = 6588/(1+z) \approx 4500 \text{ \AA}$) due to metal line blanketing. Line blanketing by Fe II and Ti II, which suppresses flux blueward of $\sim 4000 \text{ \AA}$, was observed for Type Ib SN 1999dn (Branch et al. 2002; Deng et al. 2000; Cano et al. 2014). Flux suppression due to several iron-group elements was also observed for SN 2006aj (Sollerman et al. 2006), while metal line blanketing was also suggested by Bloom et al. (1999) to explain the red colour of the SN bump of GRB 980326.

For each SN we used the luminosity factor averaged over all available filters to estimate the amount of nickel nucleosyn-

¹⁰ We note that a non-zero rest-frame extinction could partially explain this flux decrement. In our second energy spectrum we find $A_{V,\text{rest}} = 0.21^{+0.28}_{-0.21} \text{ mag}$, which we argue implies zero rest-frame extinction. However, we acknowledge that a non-zero extinction is still compatible with our result, which would partially alleviate this issue.

Table 2. GRB-SNe: Observational and Physical Properties (*UBVRIJH* rest-frame wavelength range)

GRB	SN	Filter (obs)	k	s	$T_{\text{peak,obs}}$ (d)	$M_{\text{peak,obs}}$	M_{Ni} (M_{\odot})	M_{ej} (M_{\odot}) [†]	E_K (10^{51} erg) [†]
120729A	-	<i>r</i>	1.29 ± 0.19	1.0 (fixed)	-	-18.96 ± 0.15	-	-	-
120729A	-	<i>i</i>	0.76 ± 0.11	1.0 (fixed)	-	-19.29 ± 0.15	-	-	-
120729A	-	average	1.02 ± 0.26	1.0 (fixed)	-	-	0.42 ± 0.11	-	-
130215A	2013ez	<i>r</i>	$0.6 - 0.75$	1.0 (fixed)	-	-18.7 to -19.0	-	-	-
130215A	2013ez	<i>i</i>	$0.6 - 0.75$	1.0 (fixed)	-	-19.0 to -19.3	-	-	-
130215A	2013ez	average	$0.6 - 0.75$	1.0 (fixed)	-	-	$0.25 - 0.30$	-	-
130831A	2013fu	<i>r</i>	0.65 ± 0.03	0.84 ± 0.03	18.60 ± 0.67	-18.89 ± 0.05	-	-	-
130831A	2013fu	<i>i</i>	1.04 ± 0.05	0.81 ± 0.03	18.53 ± 0.69	-19.56 ± 0.05	-	-	-
130831A	2013fu	<i>z</i>	$1.02 \pm 0.19^{\ddagger}$	0.81 (fixed)	≈ 19.1	-19.47 ± 0.19	-	-	-
130831A	2013fu	average (<i>r</i> & <i>i</i>)	0.85 ± 0.20	0.83 ± 0.02	-	-	0.30 ± 0.07	$4.71^{+0.79}_{-0.59}$	$18.7^{+9.0}_{-6.2}$
		SN type	M_{Ni} (M_{\odot})	M_{ej} (M_{\odot})	E_K (10^{51} erg)	median values from C13.			
		Ib	0.16 ($\sigma = 0.22$)	3.89 ($\sigma = 2.77$)	2.3 ($\sigma = 2.6$)				
		Ic	0.19 ($\sigma = 0.19$)	3.40 ($\sigma = 4.51$)	2.2 ($\sigma = 3.3$)				
		Ibc	0.18 ($\sigma = 0.21$)	3.56 ($\sigma = 3.51$)	2.2 ($\sigma = 2.8$)				
		Ic-BL	0.26 ($\sigma = 0.33$)	3.90 ($\sigma = 3.44$)	10.9 ($\sigma = 8.9$)				
		GRB/XRF	0.34 ($\sigma = 0.24$)	5.91 ($\sigma = 3.87$)	22.3 ($\sigma = 15.2$)				

[†] Ejecta mass and kinetic energy are calculated using the average peak photospheric velocity of $v_{\text{ph}} \approx (20 \pm 2.5) \times 10^3 \text{ km s}^{-1}$ determined for a sample of seven GRB-SNe in C13.

[‡] *z* observations of GRB 130831A are not host-subtracted, and are not considered when calculating the average luminosity and stretch factors.

thesized during the explosion, while for SN 2013fu we used the well-sampled LCs to estimate the ejecta mass and kinetic energy of the SN. We do not have knowledge of the photospheric velocity of the SN at peak light, which is a necessary ingredient of the analytical model of Arnett (1982) to estimate M_{ej} and E_K . Instead we used the average photospheric velocity for a sample of GRB-SNe presented in C13. A summary of the observational and physical properties of our three GRB-SNe are presented in Table 2.

When estimating the bolometric properties of the GRB-SNe presented here, we have used the method described in C13, where a lengthy discussion of the caveats of this modelling can be found. C13 determined the median bolometric properties of a sample of 20 GRB-SNe, finding: $M_{\text{Ni}} \sim 0.3 - 0.35 M_{\odot}$, $M_{\text{ej}} \sim 6.0 M_{\odot}$ and $E_K \sim 2.0 \times 10^{52} \text{ erg}$. The nickel masses derived here agree well with the range in C13, while the ejecta mass and kinetic energy of SN 2013fu is similar to those found in C13. In terms of physical properties, the GRB-SNe in this paper are quite typical of other GRB-SNe. In contrast, the nickel masses are much higher than those seen for SNe Ibc that are not associated with GRBs. C13 derived the median nickel masses for the largest sample of SNe Ibc yet considered (32 Ibc and nine IcBL), finding: SNe Ibc: $M_{\text{Ni}} \sim 0.15 - 0.18 M_{\odot}$, and Ic-BL: $M_{\text{Ni}} \sim 0.25 M_{\odot}$. Similarly the ejecta mass and kinetic energy of SN 2013fu is larger than those of the C13 sample of SNe Ibc ($M_{\text{ej}} \sim 3.4 - 3.9 M_{\odot}$, $E_K \sim 0.2 \times 10^{52} \text{ erg}$) and Ic-BL ($M_{\text{ej}} \sim 3.9 M_{\odot}$, $E_K \sim 1.0 \times 10^{52} \text{ erg}$).

In Fig. 12, we have plotted these three GRBs in the $E_p - E_{\text{iso},\gamma}$ plane. We have used the sample of GRBs in Amati et al. (2008) and the values of $E_{\text{iso},\gamma}$ and E_p presented in the introduction of each GRB section. Also plotted is a simple power-law, where we have used an index of $\alpha = 0.57$, a normalization constant of $\Lambda = 94$, and the $2\text{-}\sigma$ confidence intervals in the Amati relation (Amati

et al. 2008). It is seen that GRBs 130831A and 120729A lie right on the fit, while GRB 130215A is within the upper bounds of the $2\text{-}\sigma$ confidence interval.

Taking a step back, SN 2013ez is now the second known GRB-SN, after SN 2002lt, for which it is more likely that it is of type Ic rather than Ic-BL. These results seem somewhat contrary to observations of all other GRB- and XRF-SNe, where unambiguous broad lines are seen in time-series spectra of the SNe. Up to now, the general consensus has been that GRB/XRF-SNe are therefore a rare subclass of SNe Ic-BL, which themselves are a rare subclass of the general SNe Ibc population (e.g. Podsiadlowski et al. 2004; Guetta & Della Valle 2007). But are SN 2013ez and SN 2002lt really non-broad-lined SNe Ic? One way to reconcile this growing heterogeneity of GRB-SNe is to check whether the timing of observations affects how we spectroscopically classify each accompanying SN. The spectrum of SN 2002lt was obtained at $t - t_0 \approx 15 \text{ d}$ (rest-frame), which is at a similar epoch as the timing of the GTC spectrum obtained of SN 2013ez. The time of peak light of SN 2013ez is not well constrained, and could have been anywhere from 5–10 days after maximum light (rest-frame). If the SN ejecta of these two events are less dense than expected, then the ejecta will become optically thin more rapidly than in the case of other GRB-SNe, and thus much more difficult to identify as a type Ic-BL. An example of this is Ic-BL SN 2002ap (Mazzali et al. 2002) that showed broad features in the pre-maximum spectra which quickly vanished after maximum. In their spectral modelling, the photospheric velocity rapidly decreased from $\sim 30,000 \text{ km s}^{-1}$ at two days after explosion, to $\sim 20,000 \text{ km s}^{-1}$ only three days later. Therefore, it is not unfathomable that SN 2013ez may have had more rapid velocities earlier on, but they disappeared by the time we obtained our spectrum. This clearly highlights the need to obtain more than one spectrum of the accompanying SN in order

to more accurately quantify the line velocities and its spectral classification.

5.2. The magnetar model

In this work we derived the initial spin period and magnetic field of a possible millisecond magnetar central engine for GRB 130215A from optical/NIR observations. We constructed a bolometric LC from contemporaneous *grizJH* observations of GRB 130215A and fit it with the model from Zhang & Mészáros (2001; see as well Rowlinson et al. 2013), finding $P_0 = 12.0$ ms, a magnetic field strength of $B = 1.1 \times 10^{15}$ G, a plateau luminosity of $L_{\text{plat}} = 6.1 \times 10^{44}$ erg s $^{-1}$ and a plateau duration (rest-frame) of $T_{\text{plat}} = 2.3 \times 10^5$ s. These values are not unrealistic, and reminiscent of those estimated for other long (GRB 060729: $P = 1.5$ – 2.0 ms and $B = 0.25$ – 3.2×10^{15} G, Xu et al. 2009; Dall’Osso et al. 2011; Lu & Zhang 2014; and GRB 130427A: $P \sim 20$ ms and $B \sim 10^{16}$ G, Bernardini et al. 2014), a sample of 10 long-duration GRBs in Lyons et al. (2010; their Table 2), and a sample of short GRBs detected by *Swift* up to 2012 (see, e.g., Table 3 in Rowlinson et al. 2013).

We must consider the limitations of our data when interpreting this result. The bolometric LC is constructed from observations of the AG+SN+host. In early epochs the AG will dominate, however in the latter two epochs at $t - t_0 = 8.1$ and 9.8 d there will be some contribution of flux coming from SN 2013ez. Moreover there will be a constant contribution from the underlying host galaxy in all epochs. However, at early times the host and SN contribute a negligible amount of flux, though as the AG fades the SN becomes the dominant source of flux, which too fades, leaving the host as the only source of emission (this happens only after 100 days or more). In our GTC images we have not detected the host to deep limits: $g > 26.2$, $r > 26.1$, $i > 25.1$. If the host had these magnitudes, it would contribute $\sim 3 - 8$, 25% flux at $t - t_0 = 9.8$, 25 d respectively. Obviously fainter magnitudes imply less host contribution.

5.3. Future Prospects

As discussed in the introduction, there are now almost a dozen spectroscopically associated GRB-SNe, though only a few have multi-band observations, with NIR observations severely lacking except in only the nearest GRB-SNe. The bolometric properties of GRB-SNe have been shown to be statistically different to those of non-GRB SNe Ibc, implying that non-GRB SNe Ibc arise from different physical scenarios than GRB-SNe. Without the possibility of directly detecting the progenitor star of a GRB-SNe, we must infer its properties indirectly via the application of advanced modelling techniques and simulations of the SN themselves. Analytical models presented in e.g. Drout et al. (2011) and C13 can go only so far (within a factor of ~ 2) in providing a clear description of the physical processes occurring during the SN, which themselves are highly dependent on the explosion mechanism, the evolutionary stage of the progenitor at the time of explosion, and the physical processes (and timescales of those processes) powering the SN. In some cases the peak light may not be powered solely by nickel decay, but additionally by shock heating of the SN ejecta (e.g. Fryer et al. 2007). This example highlights the need to employ more sophisticated modelling techniques that incorporate increasingly more of the physical processes that are occurring during the SN, as well as a more realistic description of the geometrical arrangement of the ejecta (symmetry vs. asymmetry).

As such there is still a great need for high quality optical and NIR photometry and spectra of GRB-SNe that are then modelled. These can then be used to constrain the explosion mechanism and physical properties of the progenitor via SN modelling methods, such as Monte Carlo radiative transfer (RT) simulations (e.g. Mazzali & Lucy 1993; Maeda et al. 2006; Kasen et al. 2006). Simple RT simulations such as SYN++ (Thomas et al. 2011) provide a tool to approximately ascertain the chemical properties of the material passing through the photosphere at a given moment in time. These results can then be used as input to a RT simulation. The spectra can be used in a method such as “abundance tomography”, which has been successfully used for SNe Ia (e.g. Hachinger et al. 2013; Mazzali et al. 2014) to determine the density structure and abundance stratification in the SN ejecta. Massive stars are evolved and then exploded in hydrodynamic computer simulations, with the result being SN ejecta of a specific density structure and abundance stratification that can be directly compared with observations. In this manner a focused observing strategy aimed at obtaining optical and NIR photometry and spectroscopy with 8–10-m class ground telescopes (to obtain rest-frame *BVRi* LCs, and NIR if possible), which are combined with sophisticated simulations will undoubtedly provide deeper insight into the nature of the progenitor stars of GRB-SNe.

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Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^+	m_{err}	System	Telescope
120729A	<i>g</i>	0.0053	16.60	0.09	AB	FTN
120729A	<i>g</i>	0.0079	16.89	0.08	AB	FTN
120729A	<i>g</i>	0.0114	17.27	0.18	AB	FTN
120729A	<i>g</i>	0.0160	17.70	0.08	AB	FTN
120729A	<i>g</i>	0.0226	18.02	0.10	AB	FTN
120729A	<i>g</i>	0.0354	18.35	0.10	AB	FTN
120729A	<i>g</i>	0.0564	19.16	0.08	AB	FTN
120729A	<i>g</i>	0.0788	19.56	0.08	AB	FTN
120729A	<i>g</i>	0.6564	25.14	0.67	AB	TNG
120729A	<i>g</i>	0.7683	24.00	0.11	AB	GTC
120729A	<i>g</i>	18.6095	24.44	0.06	AB	GTC
120729A	<i>g</i>	24.6329	24.58	0.12	AB	GTC
120729A	<i>g</i>	189.3991	24.44	0.15	AB	GTC
120729A	<i>r</i>	0.7629	23.56	0.08	AB	GTC
120729A	<i>r</i>	8.7604	24.11	0.11	AB	GTC
120729A	<i>r</i>	18.5914	23.93	0.09	AB	GTC
120729A	<i>r</i>	20.5592	23.97	0.09	AB	GTC
120729A	<i>r</i>	189.3824	24.05	0.07	AB	GTC
120729A	<i>i</i>	0.0068	15.47	0.06	AB	FTN
120729A	<i>i</i>	0.0102	16.13	0.06	AB	FTN
120729A	<i>i</i>	0.0142	16.44	0.06	AB	FTN
120729A	<i>i</i>	0.0202	16.88	0.07	AB	FTN
120729A	<i>i</i>	0.0282	17.06	0.06	AB	FTN
120729A	<i>i</i>	0.0348	17.23	0.07	AB	FTN
120729A	<i>i</i>	0.0427	17.55	0.07	AB	FTN
120729A	<i>i</i>	0.0474	17.65	0.09	AB	FTN
120729A	<i>i</i>	0.0515	17.87	0.08	AB	FTN
120729A	<i>i</i>	0.0573	17.98	0.07	AB	FTN
120729A	<i>i</i>	0.0654	18.24	0.07	AB	FTN
120729A	<i>i</i>	0.0719	18.32	0.09	AB	FTN
120729A	<i>i</i>	0.0799	18.52	0.08	AB	FTN
120729A	<i>i</i>	0.6433	22.78	0.10	AB	TNG
120729A	<i>i</i>	0.6447	22.52	0.35	AB	LT
120729A	<i>i</i>	0.6468	22.64	0.10	AB	TNG
120729A	<i>i</i>	0.7575	23.24	0.06	AB	GTC
120729A	<i>i</i>	18.6275	23.66	0.07	AB	GTC
120729A	<i>i</i>	24.6534	23.63	0.07	AB	GTC
120729A	<i>i</i>	189.4159	23.89	0.07	AB	GTC
120729A	<i>z</i>	0.7738	23.27	0.27	AB	GTC
120729A	<i>z</i>	189.4331	23.86	0.14	AB	GTC
120729A	<i>V</i>	0.0060	15.91	0.10	Vega	FTN
120729A	<i>V</i>	0.5864	24.21	0.47	Vega	TNG
120729A	<i>R</i>	0.0033	15.05	0.06	Vega	FTN
120729A	<i>R</i>	0.0037	15.06	0.06	Vega	FTN
120729A	<i>R</i>	0.0041	15.21	0.07	Vega	FTN
120729A	<i>R</i>	0.0090	16.02	0.07	Vega	FTN
120729A	<i>R</i>	0.0127	16.39	0.07	Vega	FTN
120729A	<i>R</i>	0.0181	16.81	0.08	Vega	FTN
120729A	<i>R</i>	0.0254	17.09	0.08	Vega	FTN
120729A	<i>R</i>	0.0327	17.26	0.07	Vega	FTN
120729A	<i>R</i>	0.0399	17.46	0.08	Vega	FTN
120729A	<i>R</i>	0.0464	17.66	0.12	Vega	FTN
120729A	<i>R</i>	0.0500	17.70	0.10	Vega	FTN
120729A	<i>R</i>	0.0552	18.00	0.08	Vega	FTN
120729A	<i>R</i>	0.0626	18.12	0.08	Vega	FTN
120729A	<i>R</i>	0.0698	18.24	0.09	Vega	FTN
120729A	<i>R</i>	0.0771	18.54	0.08	Vega	FTN
120729A	<i>R</i>	0.0865	18.79	0.07	Vega	FTN
120729A	<i>R</i>	0.5983	22.50	0.46	Vega	LT
120729A	<i>R</i>	0.6301	23.23	0.18	Vega	TNG
120729A	I_c	0.7069	>21.3	-	Vega	IAC80

Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^\dagger	m_{err}	System	Telescope
130215A	<i>g</i>	0.9783	20.84	0.06	AB	GROND
130215A	<i>g</i>	2.9563	21.73	0.21	AB	GROND
130215A	<i>g</i>	9.7970	22.76	0.23	AB	GTC
130215A	<i>g</i>	372.8488	>26.2	-	AB	GTC
130215A	<i>r</i>	0.00810	14.09	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.00823	14.17	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00836	14.03	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00849	14.10	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00863	14.06	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00869	14.09	0.02	AB	ROTSE [†]
130215A	<i>r</i>	0.00876	14.07	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00889	14.11	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.00902	14.00	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00916	14.13	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.00929	14.16	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.00953	14.22	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.00987	14.28	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.0102	14.28	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0105	14.28	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0109	14.35	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.0110	14.34	0.01	AB	ROTSE [†]
130215A	<i>r</i>	0.0112	14.32	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.0115	14.37	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.0119	14.50	0.03	AB	ROTSE [†]
130215A	<i>r</i>	0.0122	14.47	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.0125	14.61	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.0131	14.60	0.03	AB	ROTSE [†]
130215A	<i>r</i>	0.0139	14.72	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0147	14.75	0.03	AB	ROTSE [†]
130215A	<i>r</i>	0.0155	14.83	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0163	14.86	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0167	14.92	0.02	AB	ROTSE [†]
130215A	<i>r</i>	0.0171	15.06	0.11	AB	ROTSE [†]
130215A	<i>r</i>	0.0179	15.02	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0187	15.16	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0195	15.18	0.05	AB	ROTSE [†]
130215A	<i>r</i>	0.0203	15.19	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0211	15.24	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.0219	15.30	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0227	15.44	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0235	15.45	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0243	15.56	0.13	AB	ROTSE [†]
130215A	<i>r</i>	0.0247	15.50	0.03	AB	ROTSE [†]
130215A	<i>r</i>	0.0251	15.58	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.0259	15.60	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0267	15.68	0.08	AB	ROTSE [†]
130215A	<i>r</i>	0.0275	15.64	0.11	AB	ROTSE [†]
130215A	<i>r</i>	0.0283	15.70	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0291	15.73	0.07	AB	ROTSE [†]
130215A	<i>r</i>	0.0299	15.85	0.13	AB	ROTSE [†]
130215A	<i>r</i>	0.0307	15.74	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0316	15.79	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0324	15.86	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0328	15.92	0.03	AB	ROTSE [†]
130215A	<i>r</i>	0.0332	15.93	0.09	AB	ROTSE [†]
130215A	<i>r</i>	0.0341	16.05	0.09	AB	ROTSE [†]
130215A	<i>r</i>	0.0349	15.99	0.11	AB	ROTSE [†]

Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^\dagger	m_{err}	System	Telescope
130215A	<i>r</i>	0.0357	16.19	0.14	AB	ROTSE [†]
130215A	<i>r</i>	0.0366	16.12	0.16	AB	ROTSE [†]
130215A	<i>r</i>	0.0374	16.08	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0382	16.12	0.09	AB	ROTSE [†]
130215A	<i>r</i>	0.0391	16.27	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0399	16.10	0.12	AB	ROTSE [†]
130215A	<i>r</i>	0.0407	16.22	0.17	AB	ROTSE [†]
130215A	<i>r</i>	0.0412	16.21	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0416	16.20	0.09	AB	ROTSE [†]
130215A	<i>r</i>	0.0424	16.23	0.16	AB	ROTSE [†]
130215A	<i>r</i>	0.0432	16.21	0.11	AB	ROTSE [†]
130215A	<i>r</i>	0.0441	16.19	0.11	AB	ROTSE [†]
130215A	<i>r</i>	0.0449	16.18	0.10	AB	ROTSE [†]
130215A	<i>r</i>	0.0457	16.30	0.15	AB	ROTSE [†]
130215A	<i>r</i>	0.0466	16.27	0.11	AB	ROTSE [†]
130215A	<i>r</i>	0.0474	16.50	0.22	AB	ROTSE [†]
130215A	<i>r</i>	0.0482	16.53	0.15	AB	ROTSE [†]
130215A	<i>r</i>	0.0491	16.45	0.18	AB	ROTSE [†]
130215A	<i>r</i>	0.0495	16.44	0.04	AB	ROTSE [†]
130215A	<i>r</i>	0.0499	16.28	0.13	AB	ROTSE [†]
130215A	<i>r</i>	0.0507	16.37	0.13	AB	ROTSE [†]
130215A	<i>r</i>	0.0516	16.64	0.18	AB	ROTSE [†]
130215A	<i>r</i>	0.0658	16.81	0.06	AB	ROTSE [†]
130215A	<i>r</i>	0.0674	16.99	0.06	AB	RATIR
130215A	<i>r</i>	0.0710	16.84	0.14	AB	ROTSE [†]
130215A	<i>r</i>	0.0754	17.15	0.04	AB	RATIR
130215A	<i>r</i>	0.0765	17.15	0.08	AB	RATIR
130215A	<i>r</i>	0.0824	17.19	0.06	AB	RATIR
130215A	<i>r</i>	0.0833	17.22	0.06	AB	RATIR
130215A	<i>r</i>	0.0893	17.33	0.08	AB	RATIR
130215A	<i>r</i>	0.0928	17.41	0.09	AB	RATIR
130215A	<i>r</i>	0.0938	17.48	0.04	AB	RATIR
130215A	<i>r</i>	0.0952	17.42	0.03	AB	RATIR
130215A	<i>r</i>	0.0963	17.31	0.07	AB	RATIR
130215A	<i>r</i>	0.0973	17.43	0.06	AB	RATIR
130215A	<i>r</i>	0.1021	17.49	0.04	AB	RATIR
130215A	<i>r</i>	0.1056	17.52	0.03	AB	RATIR
130215A	<i>r</i>	0.1066	17.51	0.04	AB	RATIR
130215A	<i>r</i>	0.1102	17.68	0.05	AB	RATIR
130215A	<i>r</i>	0.1111	17.84	0.06	AB	RATIR
130215A	<i>r</i>	0.1129	17.58	0.02	AB	RATIR
130215A	<i>r</i>	0.1146	17.55	0.04	AB	RATIR
130215A	<i>r</i>	0.1202	17.85	0.09	AB	RATIR
130215A	<i>r</i>	0.1362	17.70	0.22	AB	ROTSE [†]
130215A	<i>r</i>	0.1426	17.57	0.19	AB	ROTSE [†]
130215A	<i>r</i>	0.1506	17.71	0.20	AB	ROTSE [†]
130215A	<i>r</i>	0.1588	17.72	0.20	AB	ROTSE [†]
130215A	<i>r</i>	0.1667	17.55	0.23	AB	ROTSE [†]
130215A	<i>r</i>	0.9783	20.42	0.04	AB	GROND
130215A	<i>r</i>	2.0713	21.08	0.13	AB	RATIR
130215A	<i>r</i>	2.9563	21.19	0.09	AB	GROND
130215A	<i>r</i>	3.1015	21.21	0.07	AB	RATIR
130215A	<i>r</i>	4.0876	21.51	0.10	AB	RATIR
130215A	<i>r</i>	5.7866	21.54	0.09	AB	NOT
130215A	<i>r</i>	8.0828	22.59	0.18	AB	RATIR
130215A	<i>r</i>	9.7887	22.40	0.10	AB	GTC
130215A	<i>r</i>	13.8012	23.38	0.07	AB	NOT
130215A	<i>r</i>	17.5687	23.81	0.42	AB	RATIR
130215A	<i>r</i>	372.8629	>26.1	-	AB	GTC

Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^+	m_{err}	System	Telescope
130215A	<i>i</i>	0.0674	16.67	0.06	AB	RATIR
130215A	<i>i</i>	0.0685	16.61	0.17	AB	RATIR
130215A	<i>i</i>	0.0730	16.79	0.10	AB	RATIR
130215A	<i>i</i>	0.0754	16.89	0.04	AB	RATIR
130215A	<i>i</i>	0.0765	16.83	0.08	AB	RATIR
130215A	<i>i</i>	0.0824	17.00	0.07	AB	RATIR
130215A	<i>i</i>	0.0833	17.06	0.07	AB	RATIR
130215A	<i>i</i>	0.0883	17.06	0.07	AB	RATIR
130215A	<i>i</i>	0.0893	17.04	0.08	AB	RATIR
130215A	<i>i</i>	0.0928	17.08	0.07	AB	RATIR
130215A	<i>i</i>	0.0938	17.22	0.03	AB	RATIR
130215A	<i>i</i>	0.0952	17.17	0.03	AB	RATIR
130215A	<i>i</i>	0.0963	17.16	0.06	AB	RATIR
130215A	<i>i</i>	0.0973	17.25	0.07	AB	RATIR
130215A	<i>i</i>	0.1021	17.26	0.04	AB	RATIR
130215A	<i>i</i>	0.1056	17.26	0.03	AB	RATIR
130215A	<i>i</i>	0.1066	17.25	0.04	AB	RATIR
130215A	<i>i</i>	0.1102	17.34	0.05	AB	RATIR
130215A	<i>i</i>	0.1111	17.29	0.05	AB	RATIR
130215A	<i>i</i>	0.1129	17.31	0.02	AB	RATIR
130215A	<i>i</i>	0.1146	17.38	0.05	AB	RATIR
130215A	<i>i</i>	0.1202	17.51	0.09	AB	RATIR
130215A	<i>i</i>	0.8175	20.04	0.05	AB	NOT
130215A	<i>i</i>	0.9783	20.14	0.06	AB	GROND
130215A	<i>i</i>	2.0713	20.98	0.14	AB	RATIR
130215A	<i>i</i>	2.9563	20.82	0.12	AB	GROND
130215A	<i>i</i>	3.1039	20.88	0.07	AB	RATIR
130215A	<i>i</i>	4.0543	20.88	0.08	AB	RATIR
130215A	<i>i</i>	4.0876	20.81	0.06	AB	RATIR
130215A	<i>i</i>	4.1206	20.68	0.07	AB	RATIR
130215A	<i>i</i>	4.7809	20.86	0.05	AB	NOT
130215A	<i>i</i>	5.7932	21.07	0.11	AB	NOT
130215A	<i>i</i>	8.0821	22.25	0.16	AB	RATIR
130215A	<i>i</i>	9.7910	22.39	0.12	AB	GTC
130215A	<i>i</i>	11.0831	22.68	0.26	AB	RATIR
130215A	<i>i</i>	13.7889	22.78	0.08	AB	NOT
130215A	<i>i</i>	17.5687	23.52	0.40	AB	RATIR
130215A	<i>i</i>	24.2781	23.63	0.30	AB	NOT
130215A	<i>i</i>	372.8358	>25.1	-	AB	GTC
130215A	<i>z</i>	0.0671	16.35	0.20	AB	RATIR
130215A	<i>z</i>	0.0785	16.58	0.08	AB	RATIR
130215A	<i>z</i>	0.0842	16.64	0.07	AB	RATIR
130215A	<i>z</i>	0.0890	16.71	0.07	AB	RATIR
130215A	<i>z</i>	0.0904	17.09	0.13	AB	RATIR
130215A	<i>z</i>	0.0938	17.07	0.10	AB	RATIR
130215A	<i>z</i>	0.0959	16.98	0.04	AB	RATIR
130215A	<i>z</i>	0.0972	17.01	0.06	AB	RATIR
130215A	<i>z</i>	0.0980	17.15	0.10	AB	RATIR
130215A	<i>z</i>	0.1029	17.02	0.05	AB	RATIR
130215A	<i>z</i>	0.1077	17.07	0.06	AB	RATIR
130215A	<i>z</i>	0.1119	17.10	0.07	AB	RATIR
130215A	<i>z</i>	0.1139	17.11	0.04	AB	RATIR
130215A	<i>z</i>	0.1212	17.18	0.11	AB	RATIR
130215A	<i>z</i>	0.8296	20.03	0.04	AB	NOT
130215A	<i>z</i>	0.9783	19.92	0.07	AB	GROND
130215A	<i>z</i>	2.0719	20.46	0.27	AB	RATIR
130215A	<i>z</i>	2.9563	20.74	0.12	AB	GROND
130215A	<i>z</i>	3.1050	20.59	0.09	AB	RATIR
130215A	<i>z</i>	4.0855	20.63	0.09	AB	RATIR
130215A	<i>z</i>	5.8026	20.72	0.09	AB	NOT
130215A	<i>z</i>	8.0838	21.62	0.22	AB	RATIR

Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^+	m_{err}	System	Telescope
130215A	<i>z</i>	9.7940	22.27	0.14	AB	GTC
130215A	<i>Y</i>	2.0719	20.30	0.25	AB	RATIR
130215A	<i>Y</i>	3.1050	20.18	0.09	AB	RATIR
130215A	<i>Y</i>	4.0855	20.33	0.12	AB	RATIR
130215A	<i>Y</i>	8.0838	21.14	0.27	AB	RATIR
130215A	<i>J</i>	0.0754	16.38	0.10	AB	RATIR
130215A	<i>J</i>	0.0842	16.29	0.10	AB	RATIR
130215A	<i>J</i>	0.0890	16.41	0.09	AB	RATIR
130215A	<i>J</i>	0.0904	16.71	0.15	AB	RATIR
130215A	<i>J</i>	0.0939	16.31	0.10	AB	RATIR
130215A	<i>J</i>	0.0941	16.65	0.04	AB	RATIR
130215A	<i>J</i>	0.0959	16.56	0.04	AB	RATIR
130215A	<i>J</i>	0.0980	16.64	0.10	AB	RATIR
130215A	<i>J</i>	0.1028	16.66	0.06	AB	RATIR
130215A	<i>J</i>	0.1077	16.97	0.08	AB	RATIR
130215A	<i>J</i>	0.1119	16.71	0.08	AB	RATIR
130215A	<i>J</i>	0.1139	16.79	0.07	AB	RATIR
130215A	<i>J</i>	0.1212	16.82	0.12	AB	RATIR
130215A	<i>J</i>	0.9786	19.39	0.15	AB	GROND
130215A	<i>J</i>	2.0720	20.47	0.41	AB	RATIR
130215A	<i>J</i>	2.9611	20.08	0.26	AB	GROND
130215A	<i>J</i>	3.1050	20.06	0.09	AB	RATIR
130215A	<i>J</i>	4.0883	20.14	0.14	AB	RATIR
130215A	<i>J</i>	8.0842	21.55	0.50	AB	RATIR
130215A	<i>J</i>	331.8415	>23.2	-	Vega	CAHA
130215A	<i>H</i>	0.0941	15.48	0.17	AB	RATIR
130215A	<i>H</i>	0.9786	18.97	0.13	AB	GROND
130215A	<i>H</i>	2.0720	19.54	0.19	AB	RATIR
130215A	<i>H</i>	2.9611	19.94	0.31	AB	GROND
130215A	<i>H</i>	3.1050	19.79	0.10	AB	RATIR
130215A	<i>H</i>	4.0883	19.78	0.13	AB	RATIR
130215A	<i>K</i>	0.9786	18.63	0.22	AB	GROND
130215A	<i>K</i>	2.9611	18.93	0.23	AB	GROND
130831A	<i>B</i>	2.0525	22.700	0.070	Vega	MAO
130831A	<i>R_c</i>	2.0411	22.200	0.090	Vega	MAO
130831A	<i>R_c</i>	3.3168	22.200	0.110	Vega	Mondy
130831A	<i>R_c</i>	4.1061	22.530	0.090	Vega	Mondy
130831A	<i>R_c</i>	4.2555	22.530	0.070	Vega	MAO
130831A	<i>R_c</i>	5.2440	22.880	0.130	Vega	CrAO
130831A	<i>R_c</i>	5.2658	22.870	0.060	Vega	MAO
130831A	<i>R_c</i>	7.0446	22.990	0.080	Vega	MAO
130831A	<i>R_c</i>	8.0586	22.900	0.060	Vega	MAO
130831A	<i>R_c</i>	9.2558	23.110	0.060	Vega	MAO
130831A	<i>R_c</i>	10.2526	22.940	0.080	Vega	MAO
130831A	<i>R_c</i>	11.2419	22.930	0.110	Vega	MAO
130831A	<i>R_c</i>	15.2202	22.870	0.080	Vega	MAO
130831A	<i>R_c</i>	16.2715	22.770	0.070	Vega	MAO
130831A	<i>R_c</i>	22.0721	22.530	0.290	Vega	MAO
130831A	<i>R_c</i>	27.0558	22.900	0.170	Vega	MAO
130831A	<i>R_c</i>	28.6928	22.910	0.094	Vega	WHT
130831A	<i>R_c</i>	30.15945	23.360	0.160	Vega	MAO
130831A	<i>R_c</i>	37.1378	23.590	0.160	Vega	MAO
130831A	<i>R_c</i>	40.60034	23.640	0.230	Vega	Mondy
130831A	<i>R_c</i>	44.5578	> 23.2	-	Vega	Mondy
130831A	<i>R_c</i>	66.3495	23.650	0.150	Vega	CrAO
130831A	<i>R_c</i>	96.0754	23.750	0.350	Vega	Mondy
130831A	<i>I_c</i>	28.6701	22.076	0.129	Vega	WHT
130831A	<i>g</i>	0.3834	19.672	0.022	AB	NOT
130831A	<i>g</i>	4.6225	23.464	0.034	AB	WHT
130831A	<i>r</i>	0.3907	19.324	0.026	AB	NOT
130831A	<i>r</i>	0.6974	20.423	0.025	AB	NOT

Table 3. List of Photometry

GRB	Filter	$t - t_0$ (d)	m^\dagger	m_{err}	System	Telescope
130831A	<i>r</i>	0.7013	20.432	0.027	AB	NOT
130831A	<i>r</i>	7.5947	22.995	0.048	AB	NOT
130831A	<i>r</i>	13.4525	23.038	0.094	AB	NOT
130831A	<i>r</i>	27.3877	23.414	0.055	AB	NOT
130831A	<i>r</i>	68.5157	23.922	0.051	AB	Gemini
130831A	<i>r</i>	156.3151	24.063	0.089	AB	NOT
130831A	<i>i</i>	0.3901	19.100	0.027	AB	NOT
130831A	<i>i</i>	1.6584	21.741	0.066	AB	NOT
130831A	<i>i</i>	3.6759	22.403	0.053	AB	NOT
130831A	<i>i</i>	4.6024	22.540	0.044	AB	WHT
130831A	<i>i</i>	13.4761	22.580	0.061	AB	NOT
130831A	<i>i</i>	22.6051	22.422	0.058	AB	NOT
130831A	<i>i</i>	27.4342	22.843	0.054	AB	NOT
130831A	<i>i</i>	68.5281	23.777	0.089	AB	Gemini
130831A	<i>i</i>	127.2841	24.235	0.097	AB	LT
130831A	<i>z</i>	0.3987	19.026	0.043	AB	NOT
130831A	<i>z</i>	1.4975	21.218	0.099	AB	LT
130831A	<i>z</i>	6.5690	22.720	0.112	AB	LT
130831A	<i>z</i>	17.5640	22.726	0.200	AB	LT

Table 3. [†] Apparent magnitudes of the AG+SN+host are not corrected for foreground or rest-frame extinction.[‡]ROTSE magnitudes are unfiltered calibrated to *r*.

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